

INTRODUCTION TO COSMOLOGY

Joe Silk

Oxford/IAP

July 2011



CURVATURE

=

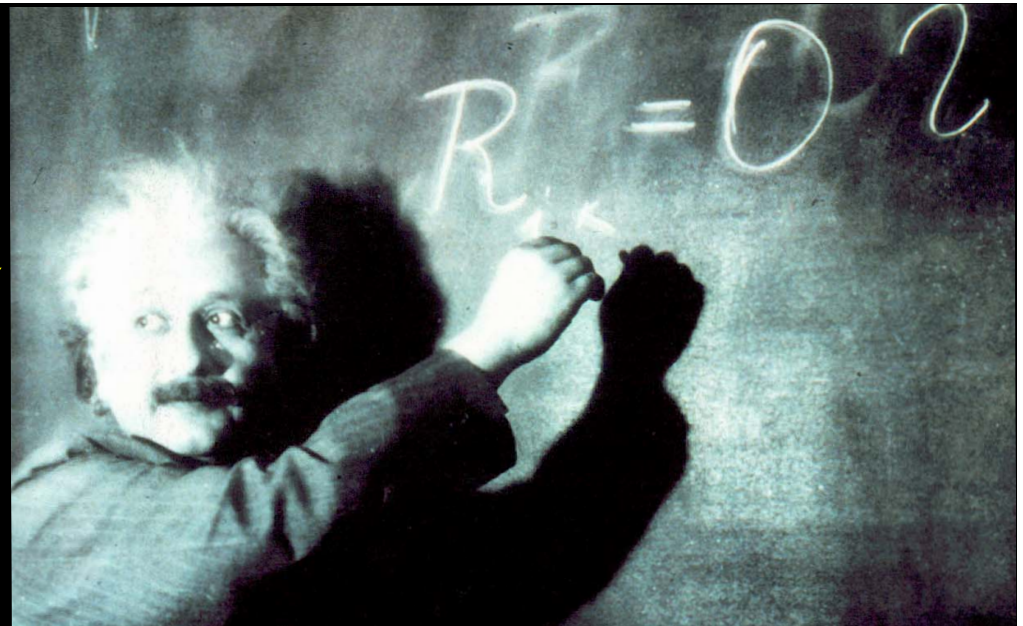
ENERGY-MOMENTUM

$G_{\mu\nu}$

=

$8\pi T_{\mu\nu}$

THE BIRTH OF DARK ENERGY



CURVATURE = ENERGY-MOMENTUM

$$G_{\mu\nu} = 8\pi T_{\mu\nu}$$

$$G_{\mu\nu} - \Lambda g_{\mu\nu} = 8\pi T_{\mu\nu}$$

$$G_{\mu\nu} = 8\pi T_{\mu\nu} + \Lambda g_{\mu\nu}$$

COSMOLOGICAL CONSTANT OR DARK ENERGY?

Field equation: $R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R - \Lambda g_{\mu\nu} = 8\pi G T_{\mu\nu}$

Perfect fluid stress tensor: $T^{\mu}_{\nu} = \text{diag}(\rho, -p, -p, -p)$

$$\tilde{T}^{\mu}_{\nu} = \text{diag}(\rho_{\Lambda}, \rho_{\Lambda}, \rho_{\Lambda}, \rho_{\Lambda}) \quad \rho_{\Lambda} = \Lambda/8\pi G$$

$$R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R = 8\pi G (T_{\mu\nu} + \tilde{T}_{\mu\nu})$$

In general, $p=w\rho$ with $w=-1$ for the cosmological constant Λ

Observational cosmology

$$\dot{R}^2 - \frac{8\pi G}{3} \rho R^2 = -kc^2$$

$$\Omega \equiv \frac{\rho}{\rho_c} = \frac{8\pi G\rho}{3H^2}$$

$$H^2(a) = H_0^2 [\Omega_v + \Omega_m a^{-3} + \Omega_r a^{-4} - (\Omega_{\text{total}} - 1)a^{-2}]$$

Distance measurement

$$D(z) = \int_0^z \frac{c}{H(z')} dz'$$

Density fluctuation growth

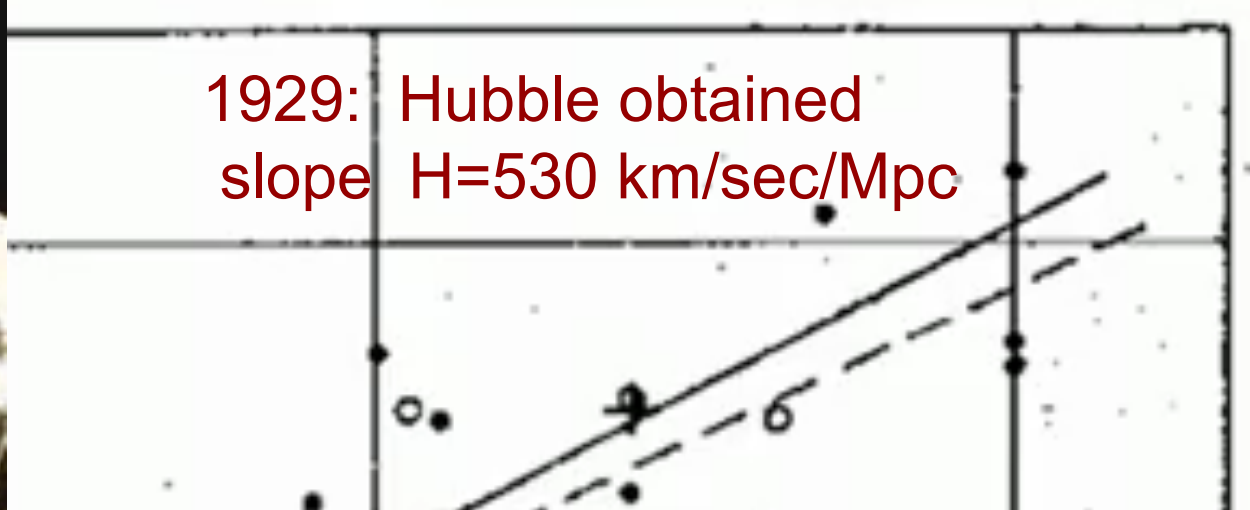
$$\delta \propto f(\Omega)a$$

$$\ddot{\delta} + 2\frac{\dot{a}}{a}\dot{\delta} = \delta \left(4\pi G\rho_m - c_s^2 k^2 / a^2 \right)$$

$$f(\Omega) \simeq \frac{5}{2}\Omega_m \left[\Omega_m^{4/7} - \Omega_v + \left(1 + \frac{1}{2}\Omega_m \right) \left(1 + \frac{1}{70}\Omega_v \right) \right]^{-1}$$



1929: Hubble obtained slope $H=530 \text{ km/sec/Mpc}$



velocity
y

distance

FIGURE 1



Alexander Friedmann:
1924

1930



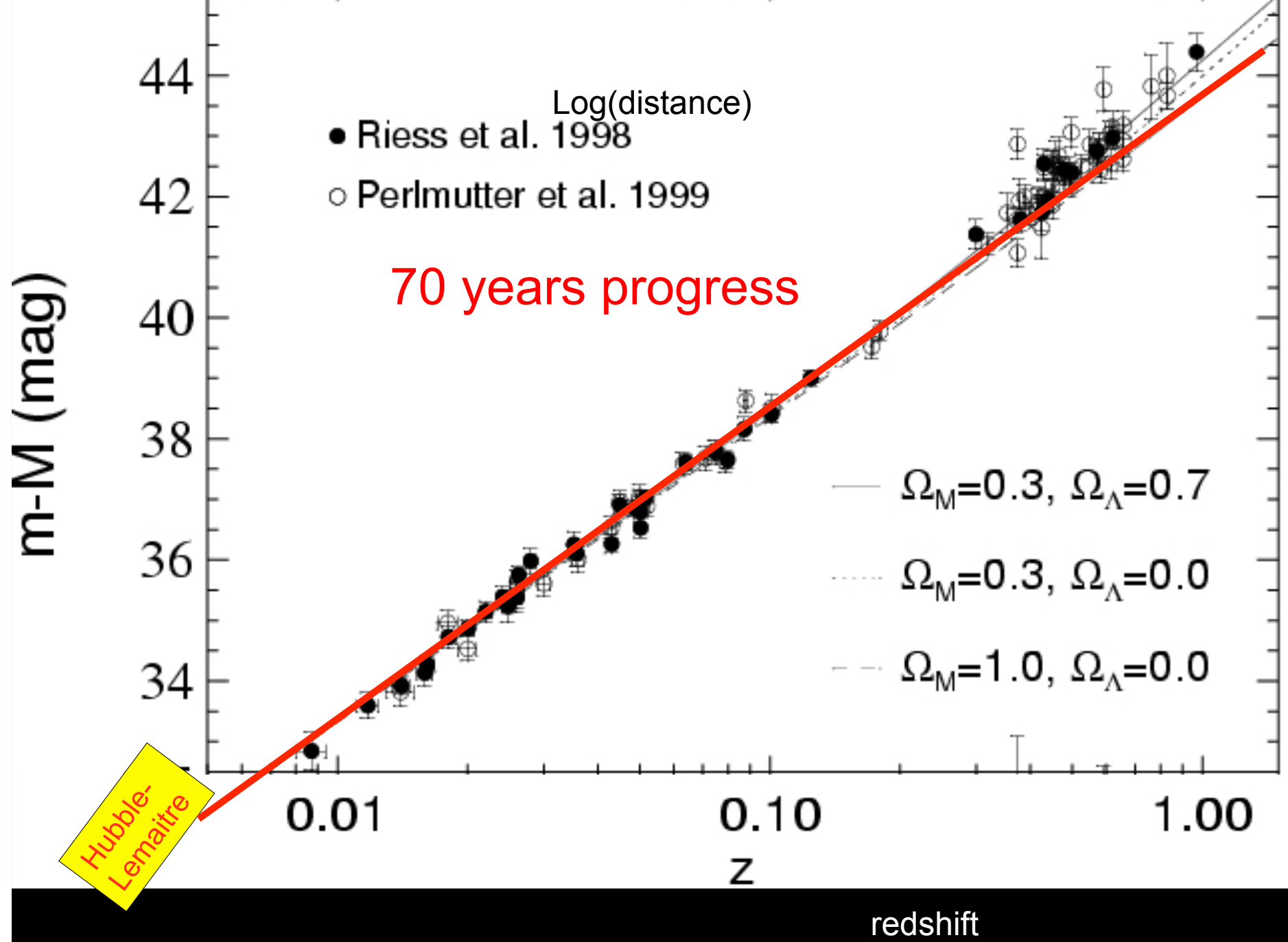
2500



1927: Georges Lemaitre obtained slope $H=625 \text{ km/sec/Mpc}$

RV (km/s)





1949

nuclear physicists plunge in

RALPH ALPHER
GEORGE GAMOW

predicted origin of light elements
& fossil radiation



The Origin of Chemical Elements

R. A. ALPHER*

*Applied Physics Laboratory, The Johns Hopkins University,
Silver Spring, Maryland*

AND

H. BETHE

Cornell University, Ithaca, New York

AND

G. GAMOW

The George Washington University, Washington, D. C.

February 18, 1948

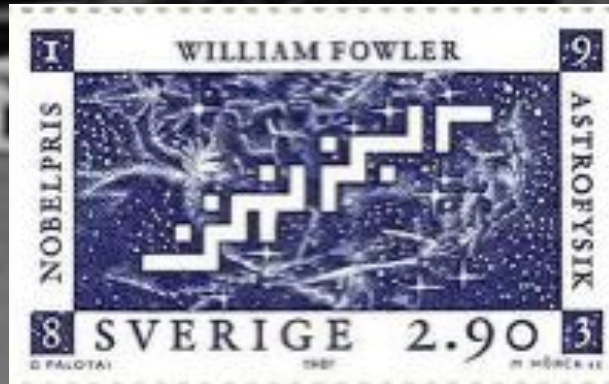
half
an
hour
of
creation...

PHYSICS TODAY
AUGUST 1950

What happened to matter in the first one thousand

At the present epoch, in which the density of matter in the universe is about 10^{-27} g/cm³, and if temperature is only about 3°K, the density of radiation (according to the Stefan-Boltzmann law) is 10^{-13} g/cm³. Thus even now the mass-density of radiation (if listed in mass-energy equivalence law) is only one twenty times smaller than that of matter.

by George Gamow



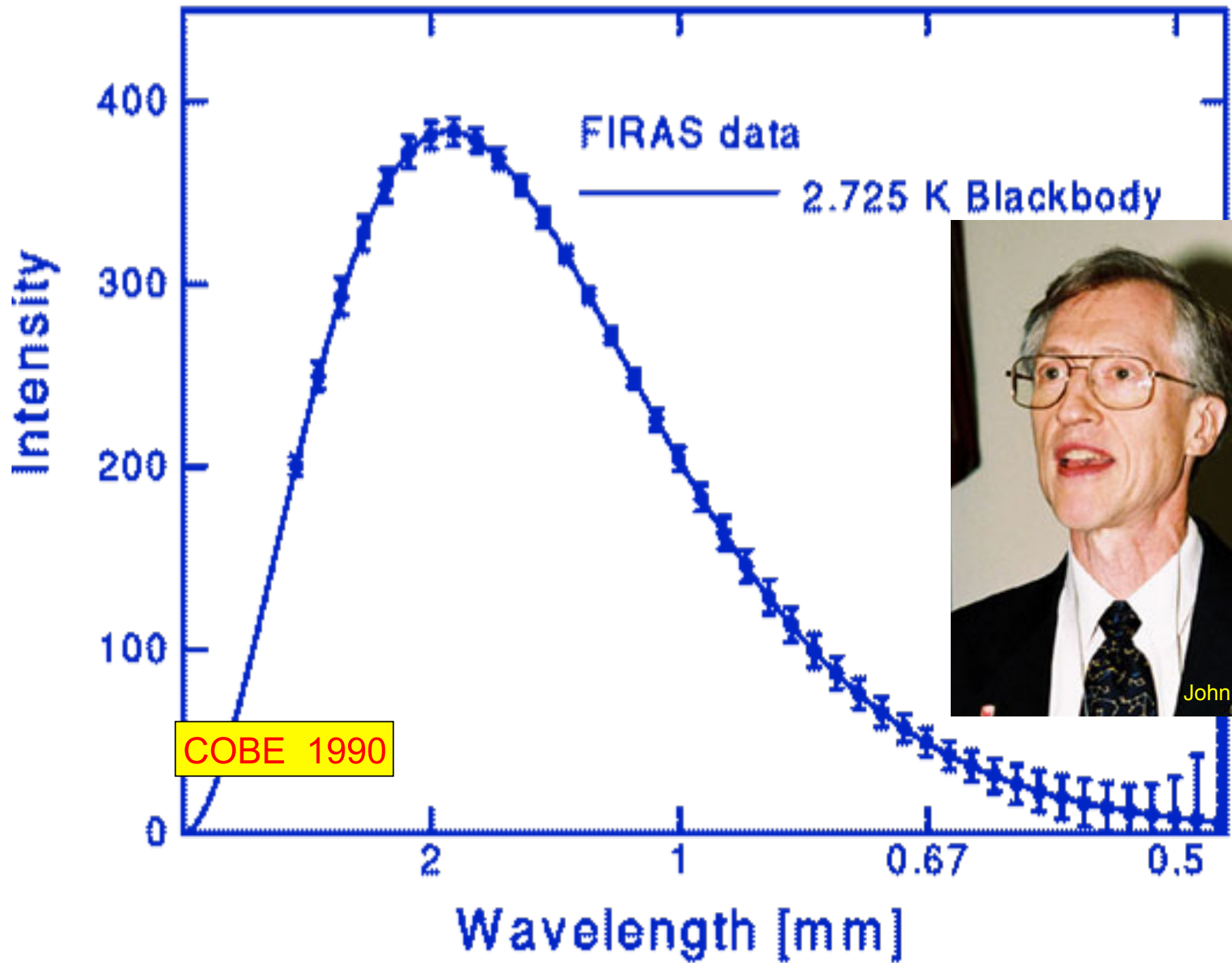
idge, Burbidge, Fowler, Ho

FROM ITS ASHES OUR SOLAR SYSTEM FORMED



1964

RELIC RADIATION DISCOVERED BY ARNO
PENZIAS AND ROBERT WILSON

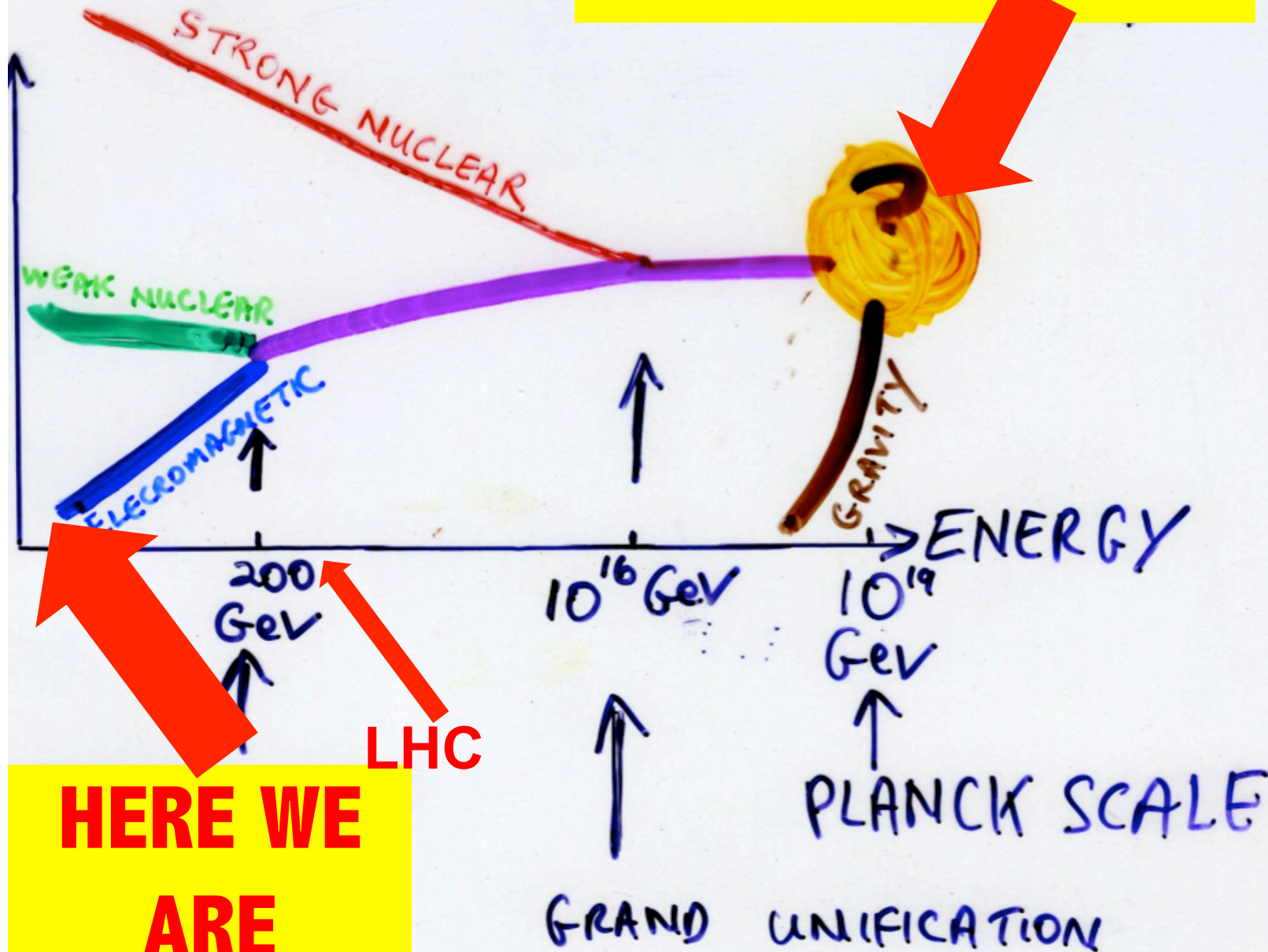


John Mather

1980

particle physicists plunge in

THE BEGINNING



FROM EXPANSION TO INFLATION (1980)



Andrei Linde



Trying to describe the size of the Big Bang

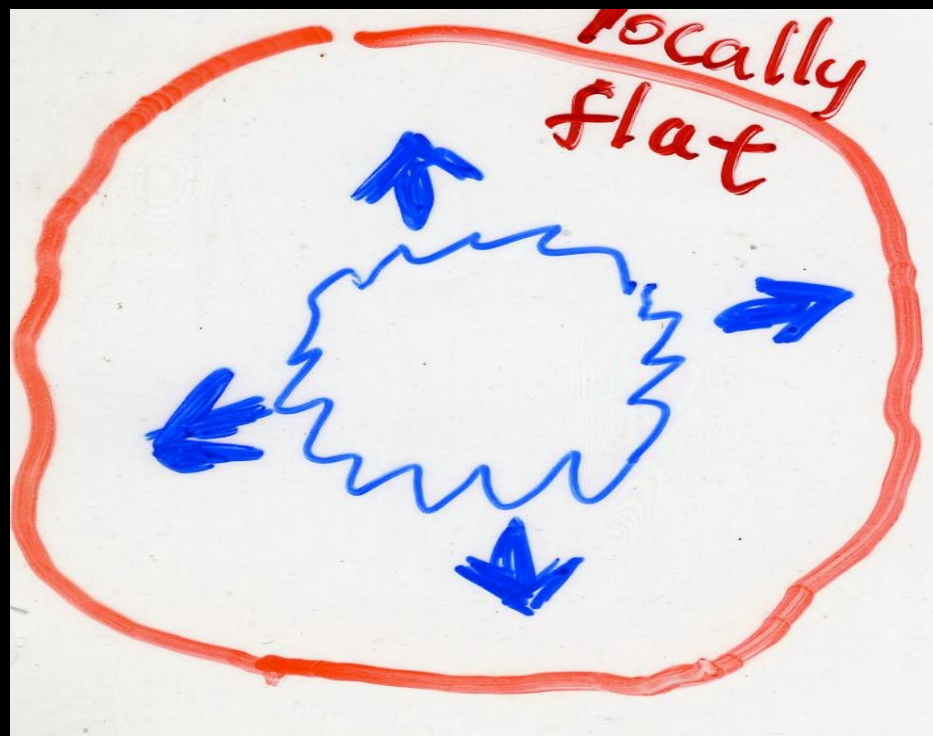


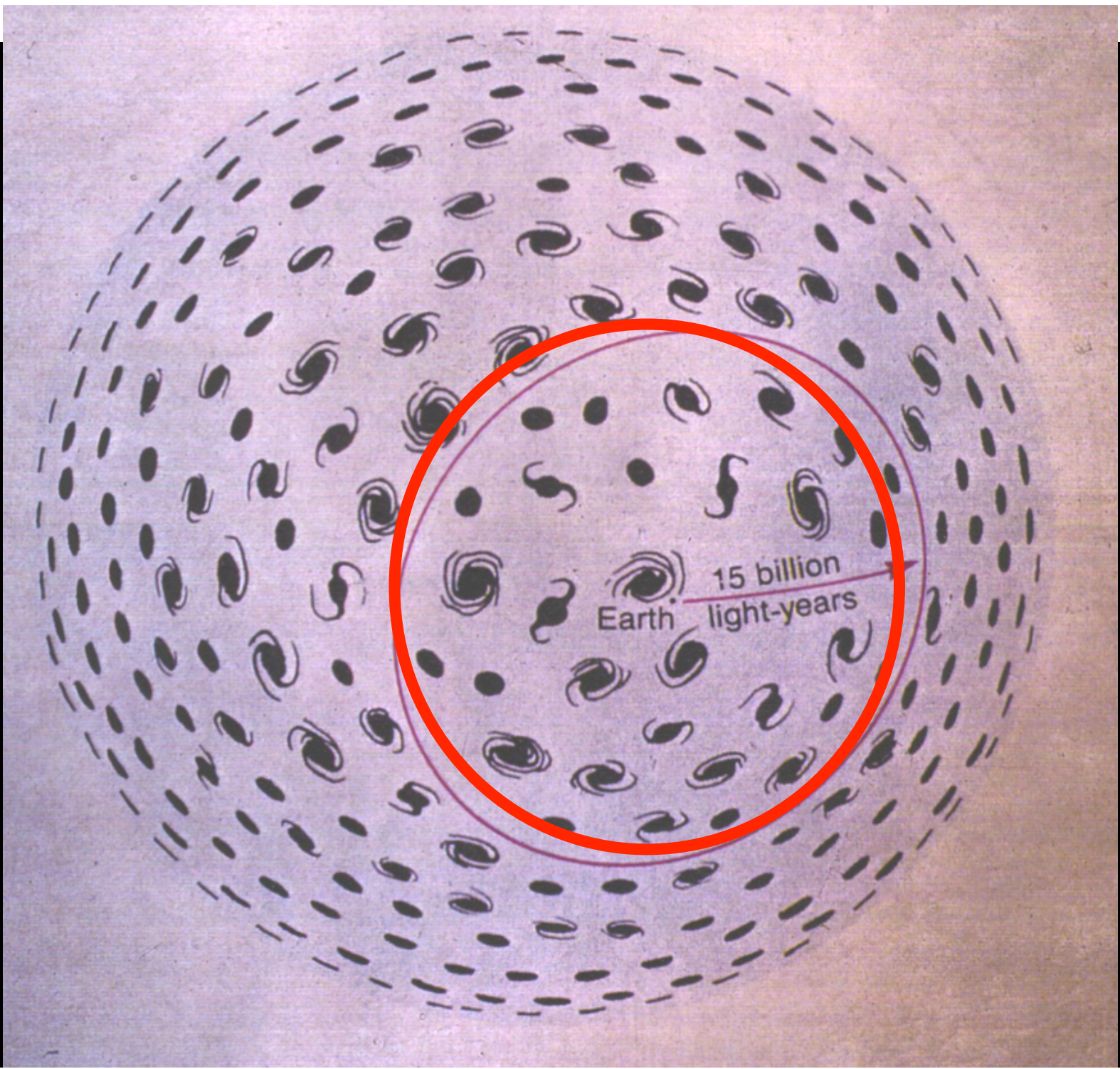
Alan Guth

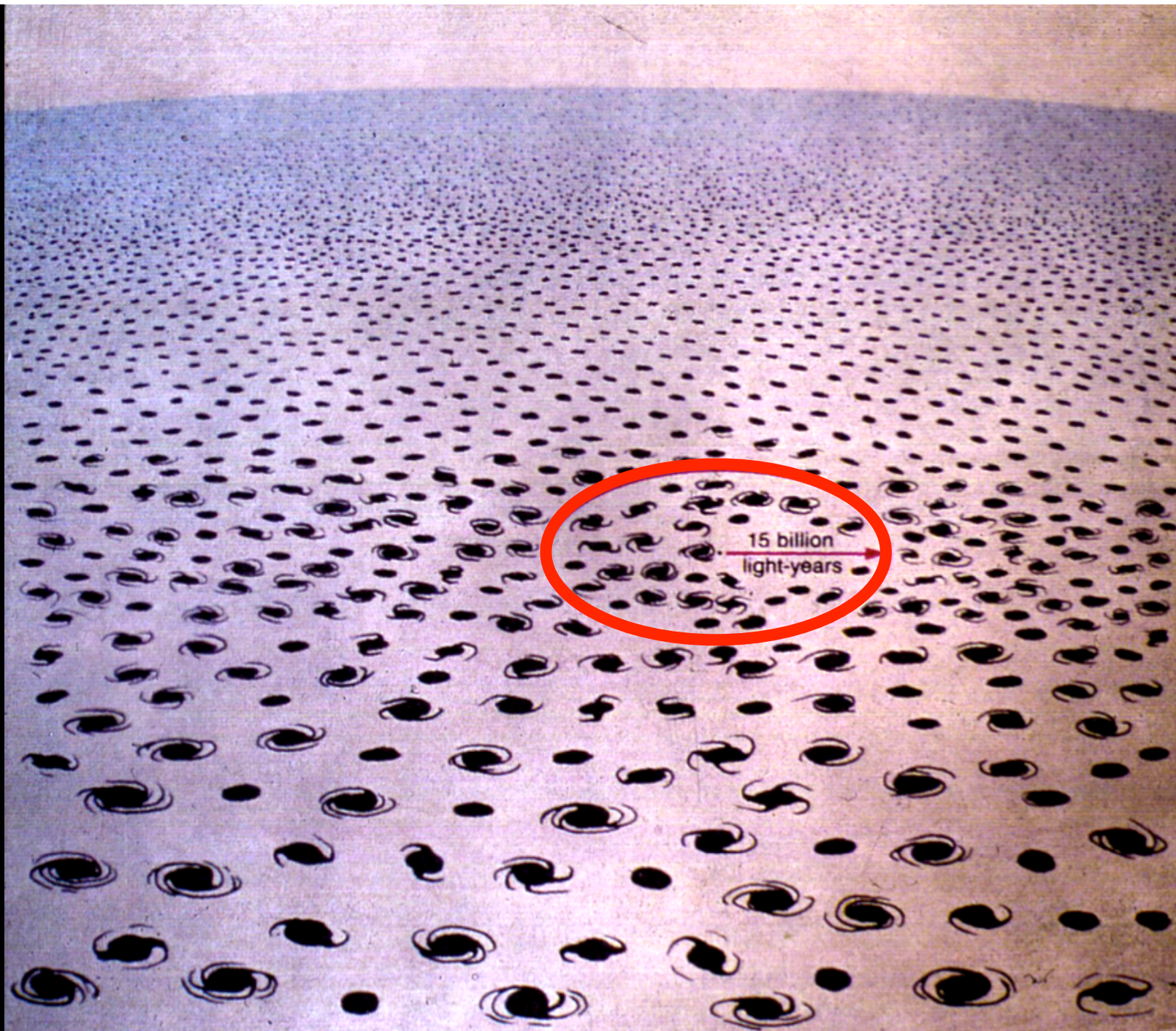
INFLATION

at 10^{-35} second

Predicts space is flat and large









THE UNIVERSE BEFORE
THE BIG BANG
(ACTUAL SIZE)

sketchy

Naive expectation:

in a static medium: $t_{\text{inst}} \sim 1/\sqrt{G\rho}$ and $\frac{\delta\rho}{\bar{\rho}} \sim e^{t/t_{\text{inst}}}$

dispersion relation: $\omega^2 = k^2 v_s^2 - 4\pi G\rho$ (Jeans instability)

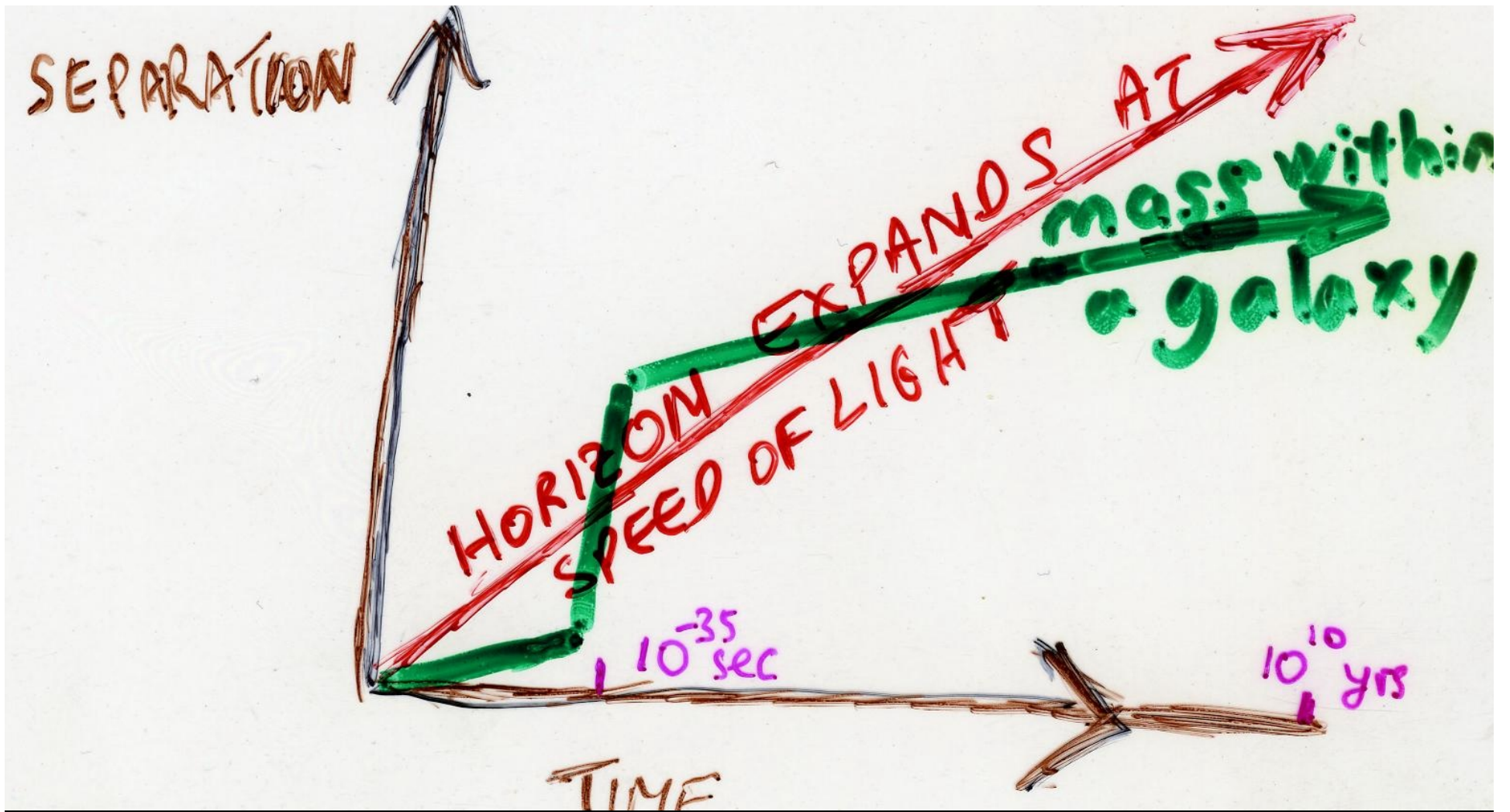
for $\delta\rho \sim \exp(i\omega t - ikx)$

with $\rho \propto t^{-2}$, expect power-law growth rate

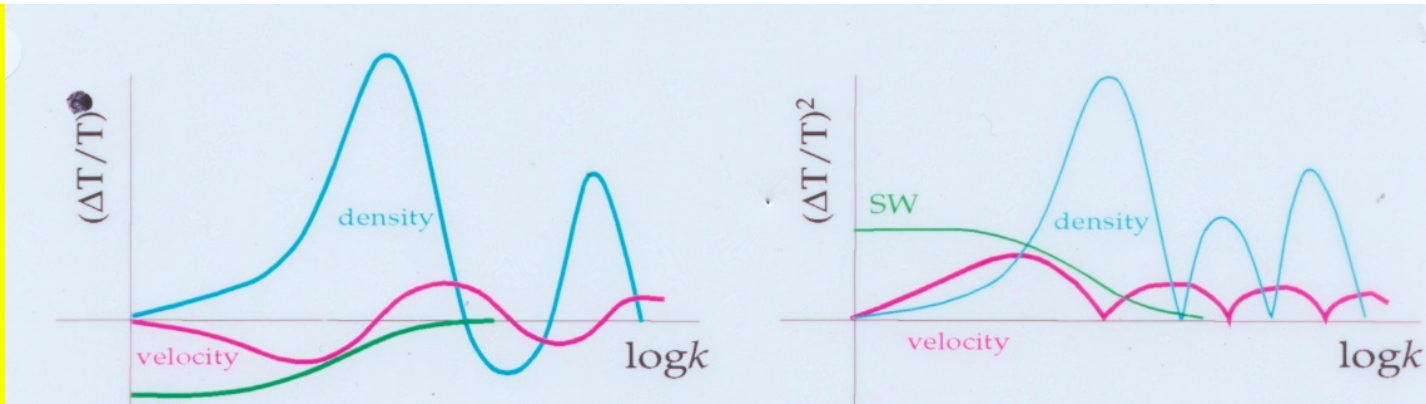
**Need initial density fluctuations
in order to form structure!**

$$\left| \frac{\delta T}{T} \right| = \left| \underbrace{\frac{1}{3} \frac{\delta\rho}{\rho} + \delta\phi}_{-\frac{1}{3} \delta\phi} - \hat{\mathbf{i}} \cdot \mathbf{v} \right|$$

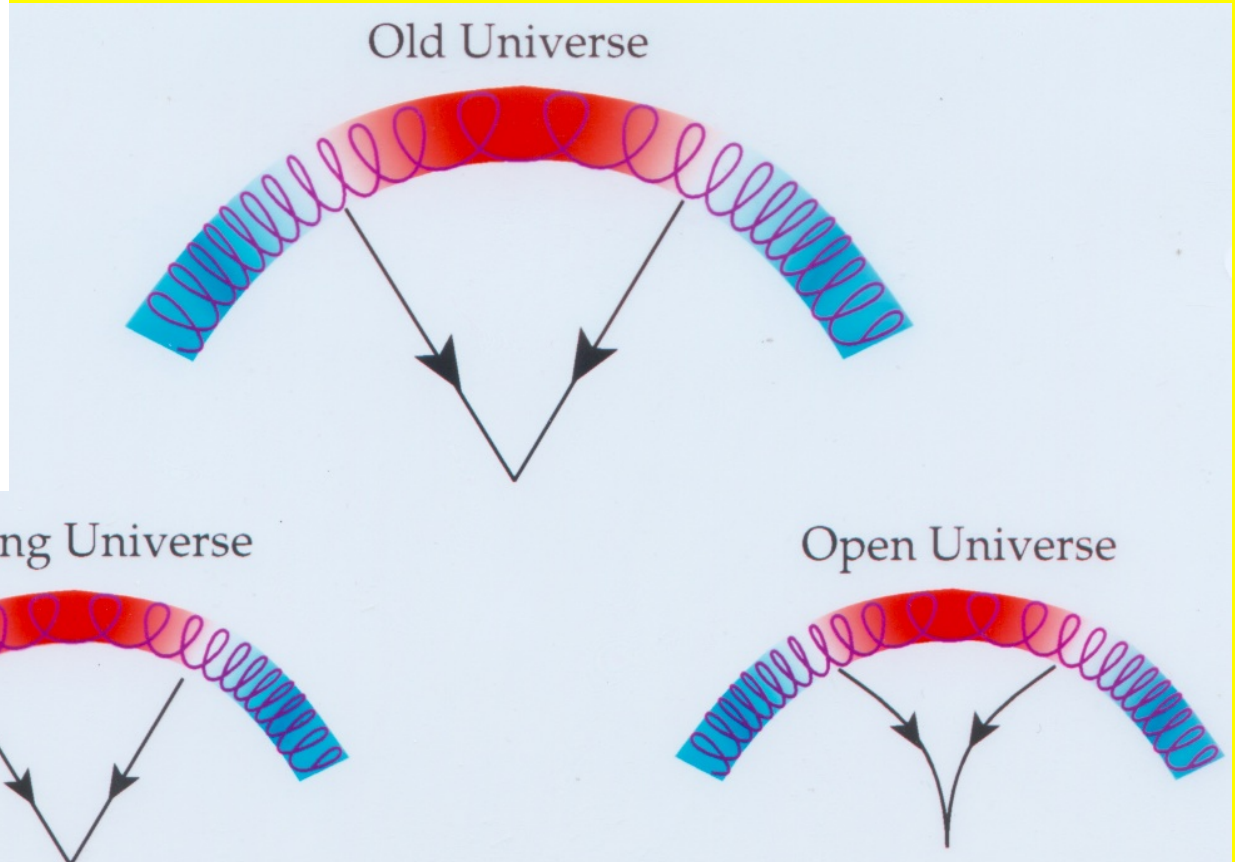
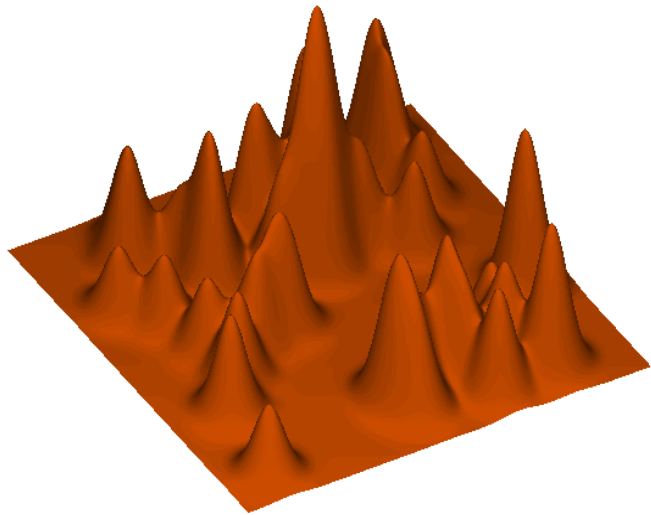
intrinsic potential doppler



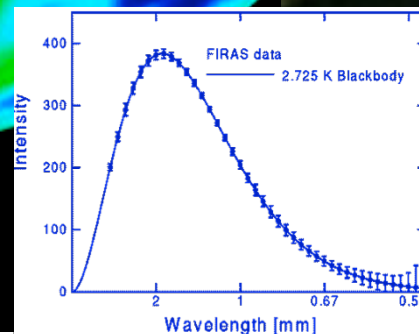
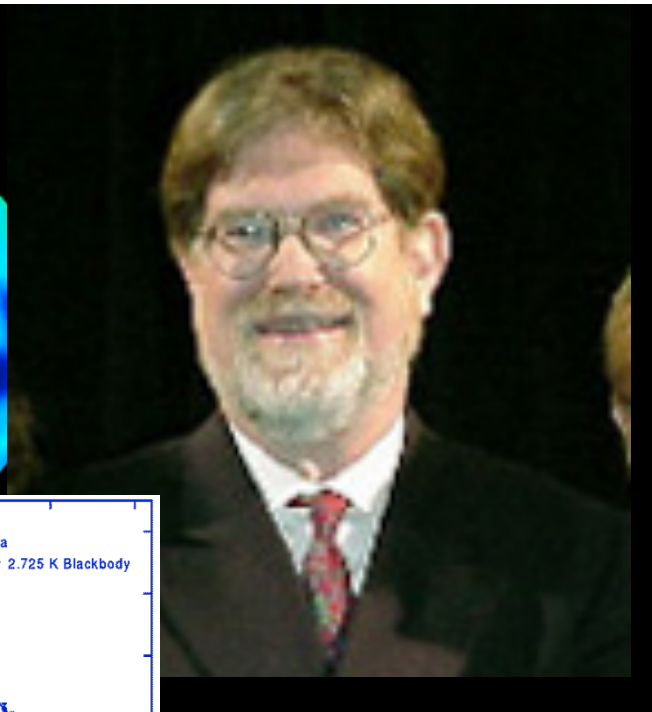
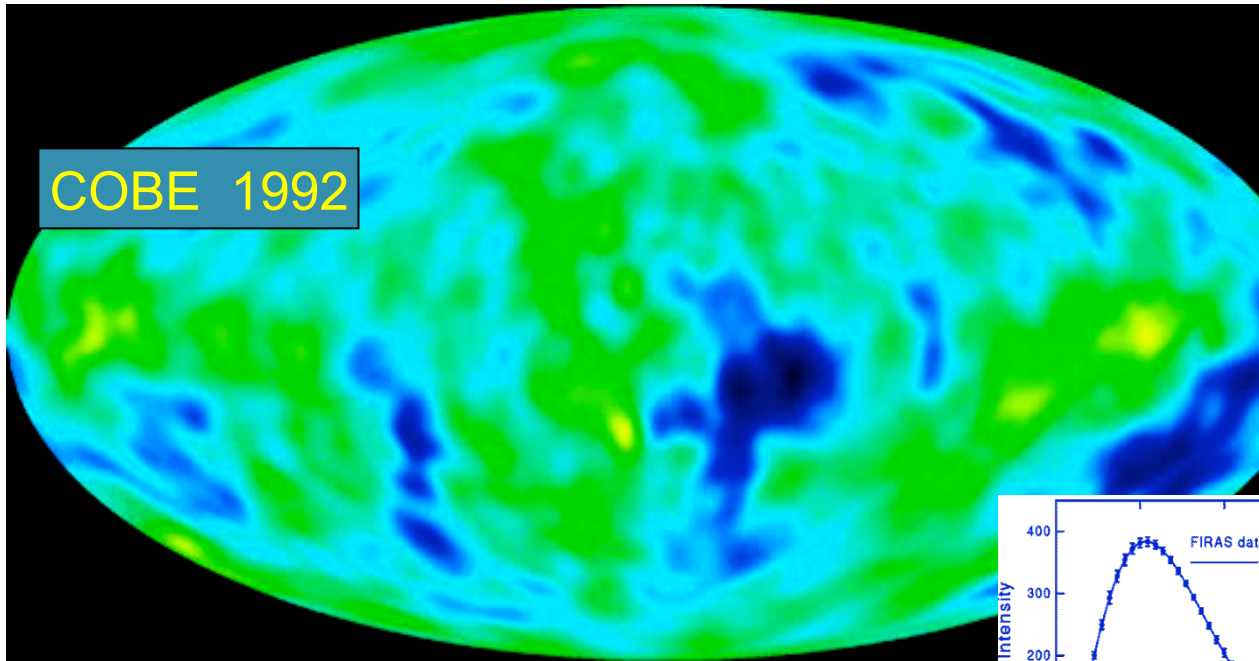
Inflation predicts thermal fluctuations



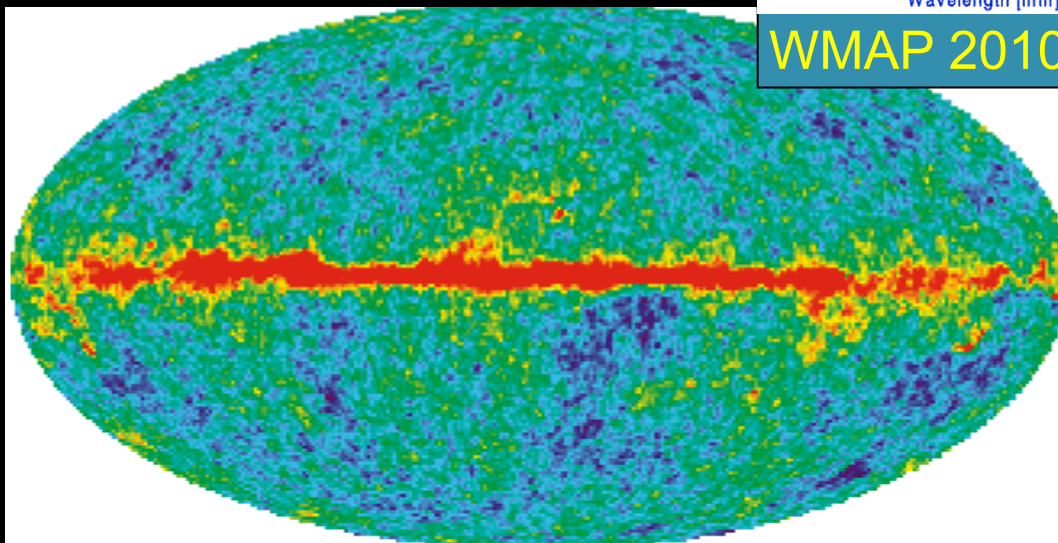
measure space curvature: dark matter plus dark energy once we know the age of the universe



COBE 1992



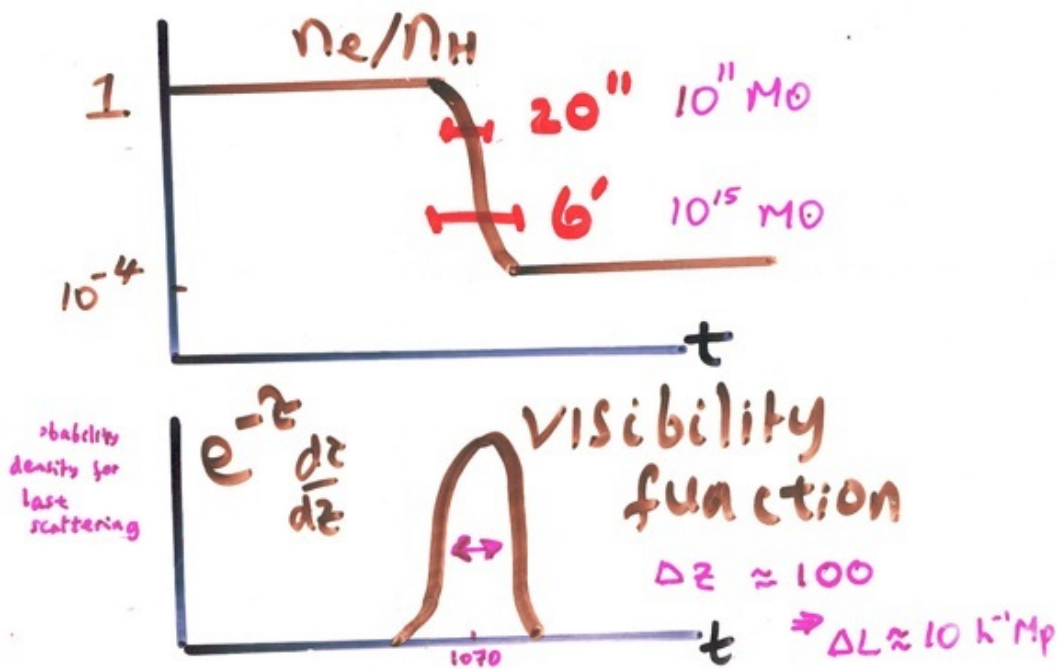
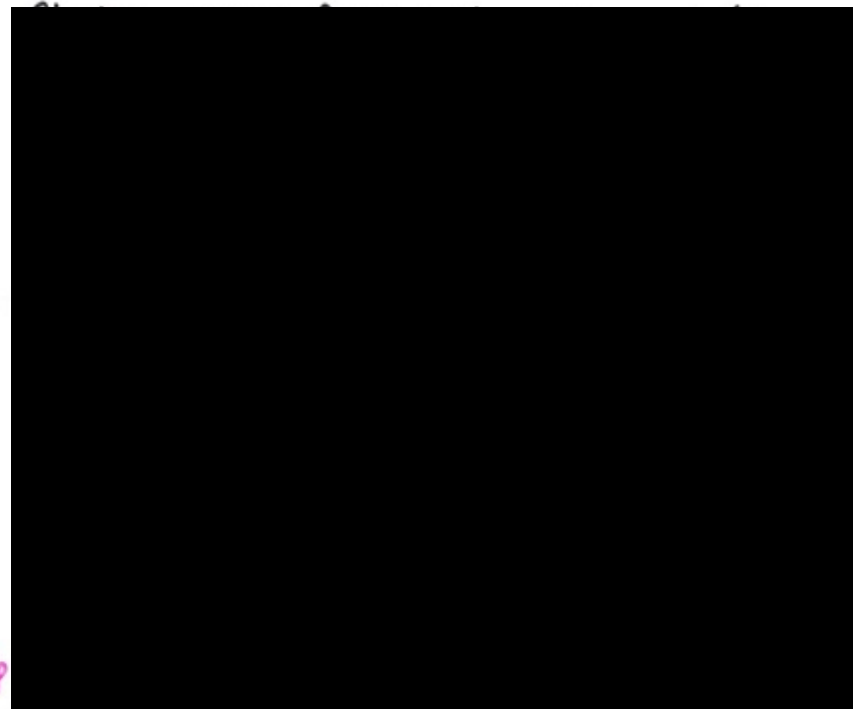
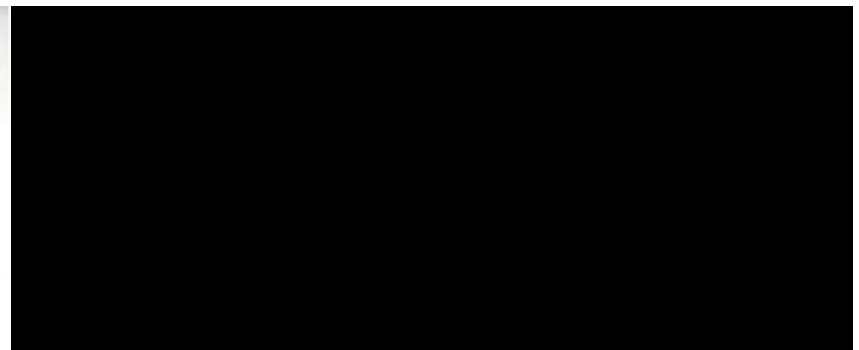
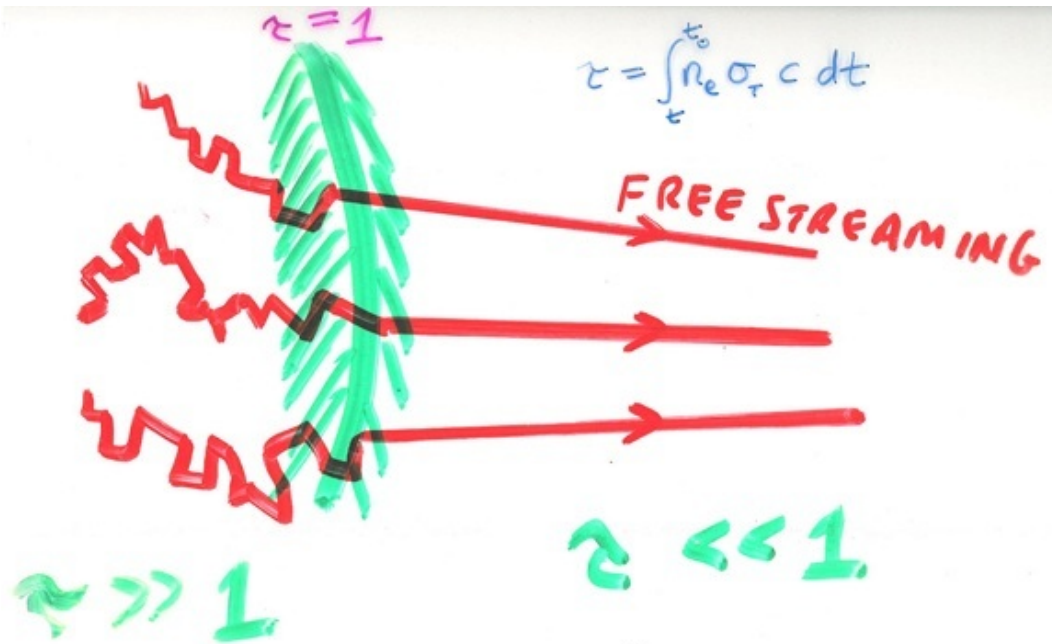
WMAP 2010



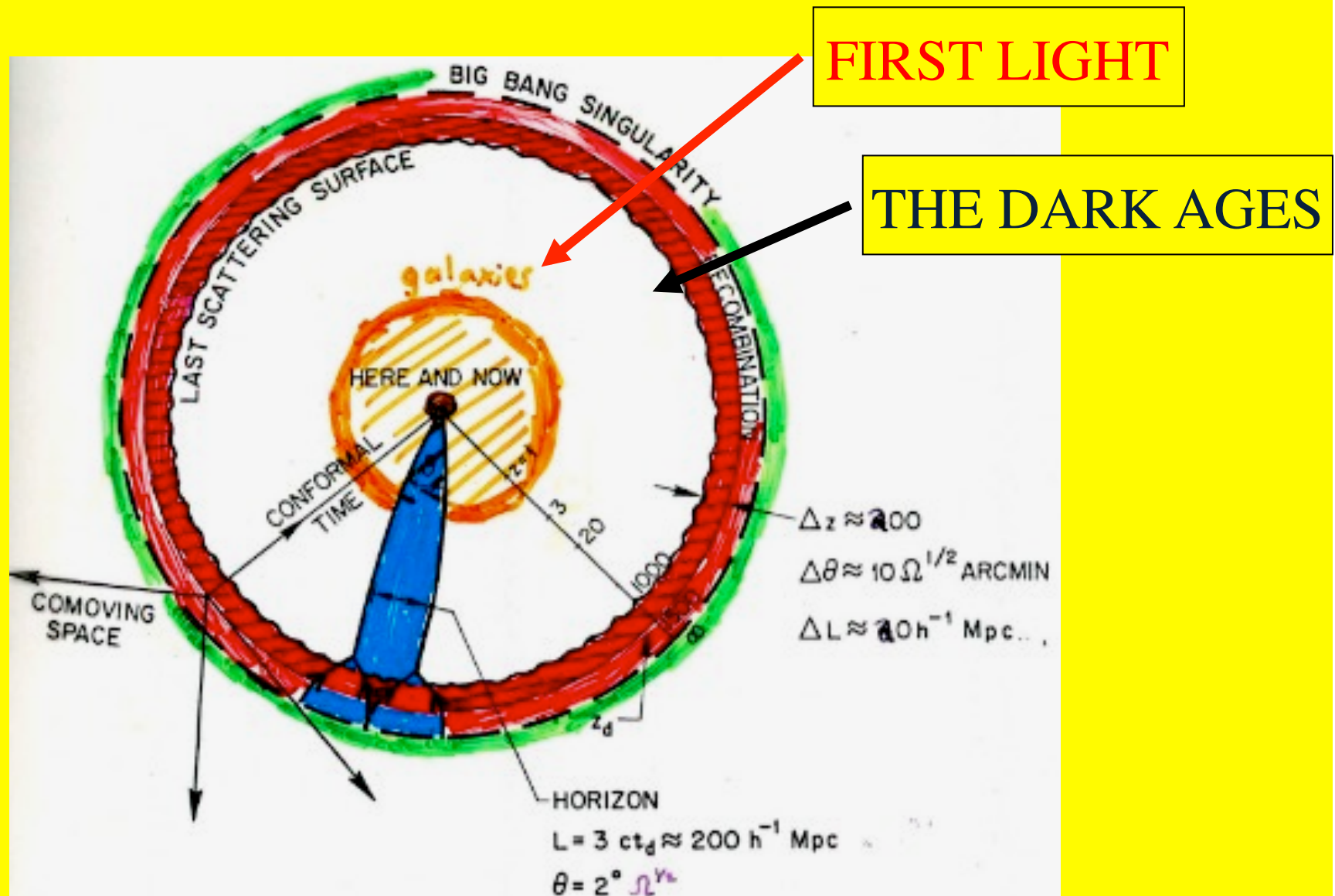
-300 μK  300 μK



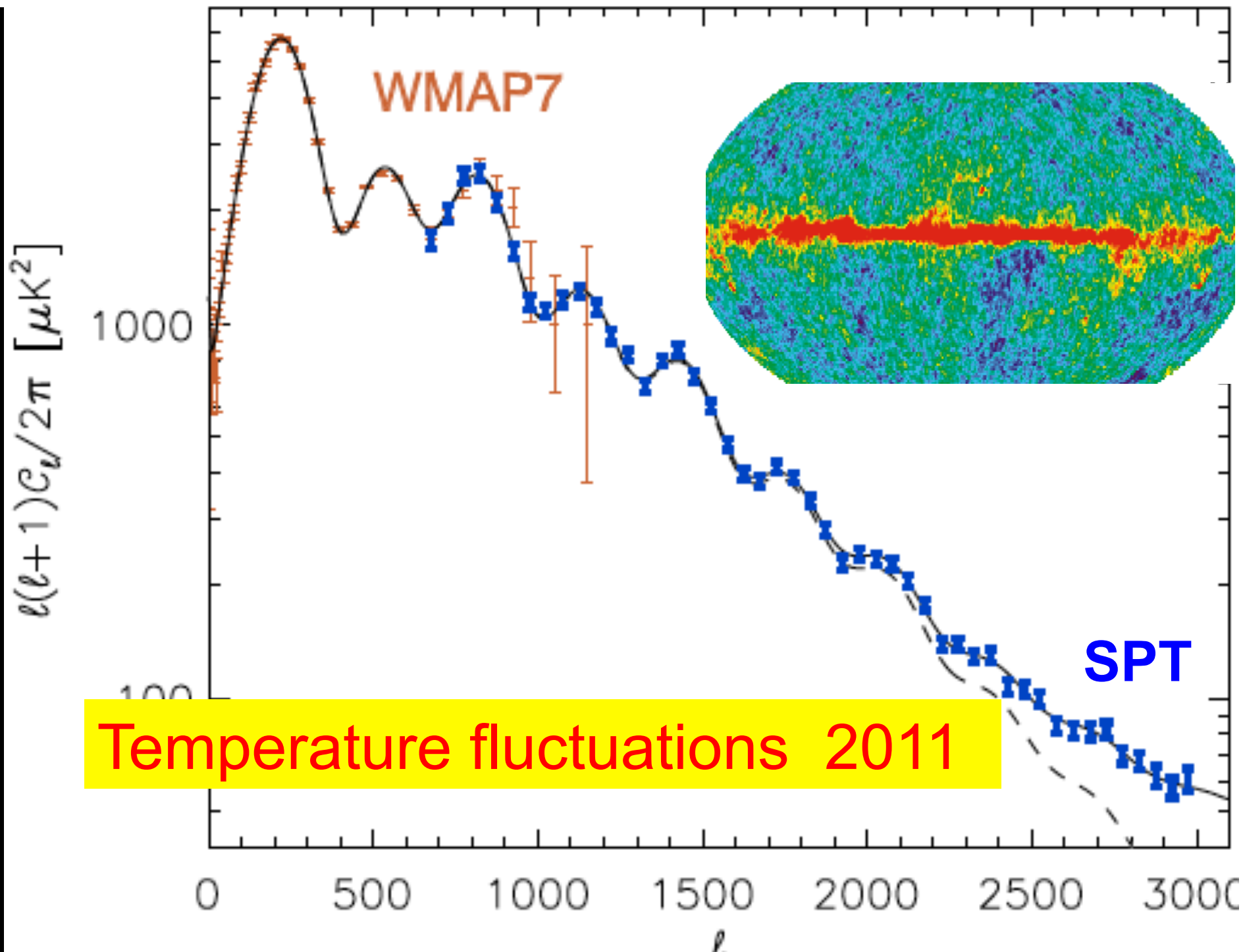
George Smoot John Mather
STOCKHOLM 2006



A SPACE-TIME DIAGRAM OF THE UNIVERSE

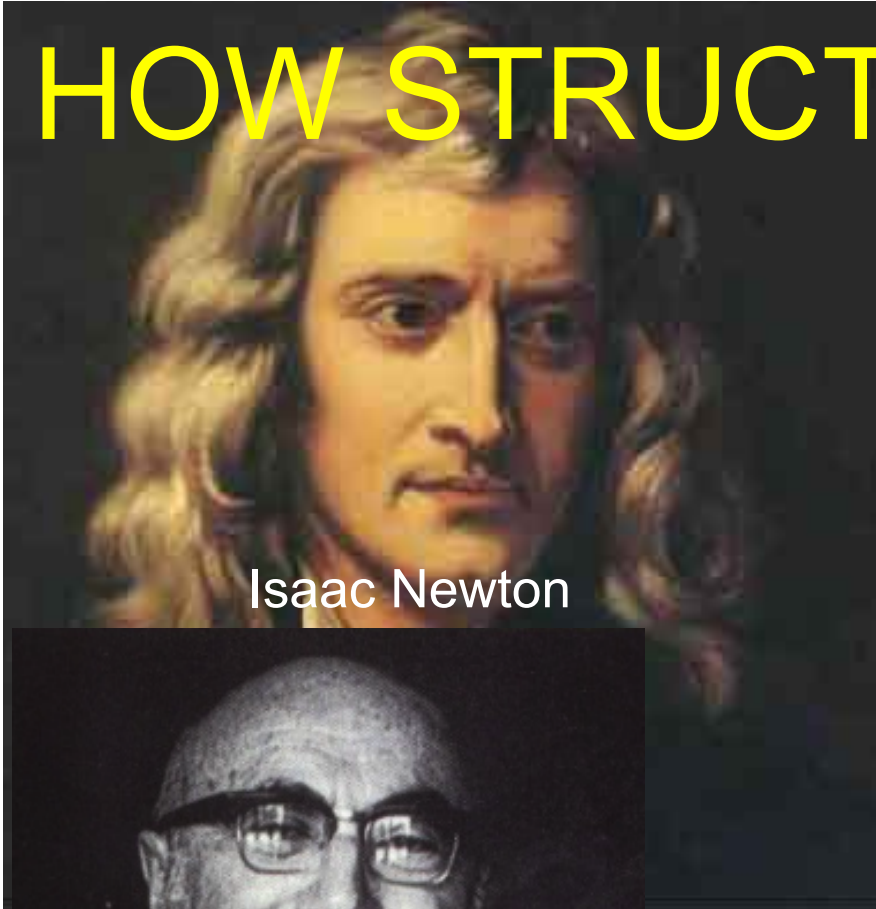


ON ANGULAR SCALES ABOVE A DEGREE, WE ARE VIEWING QUANTUM FLUCTUATIONS IN THE SKY!

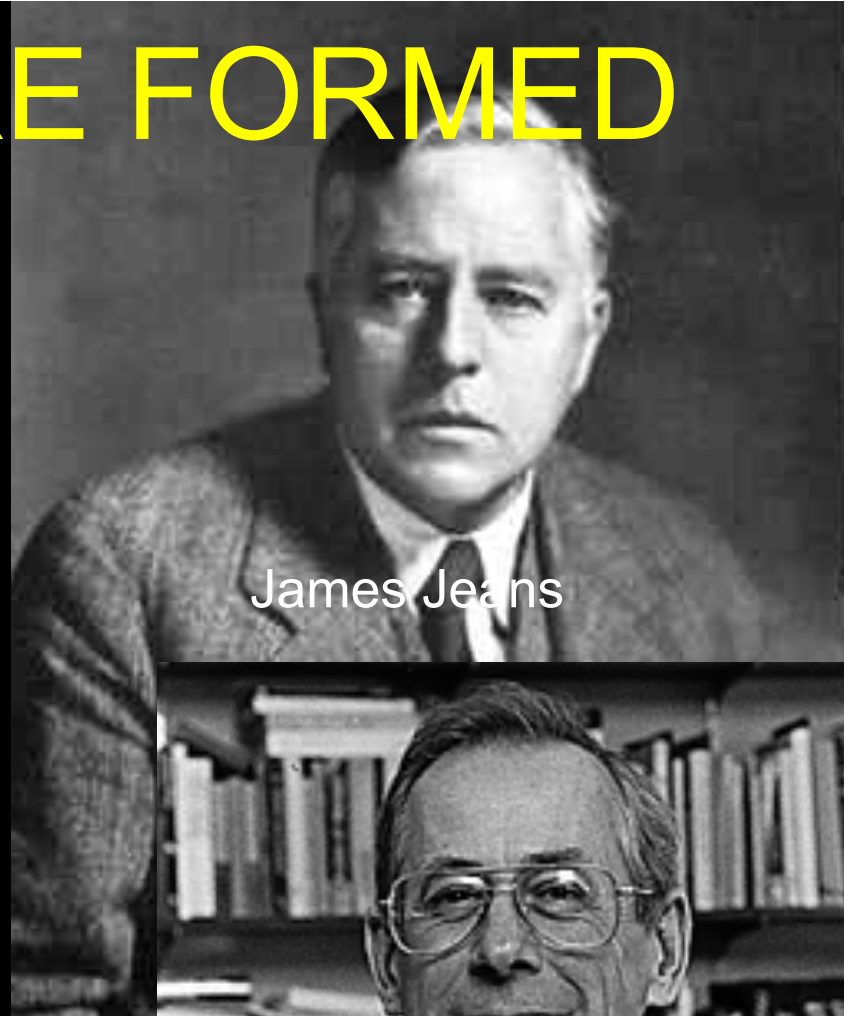


Temperature fluctuations 2011

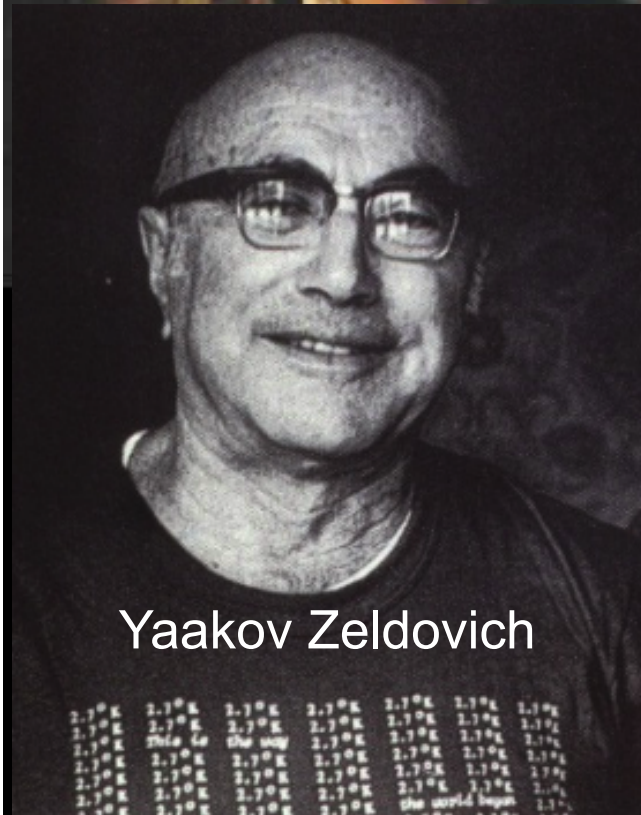
HOW STRUCTURE FORMED



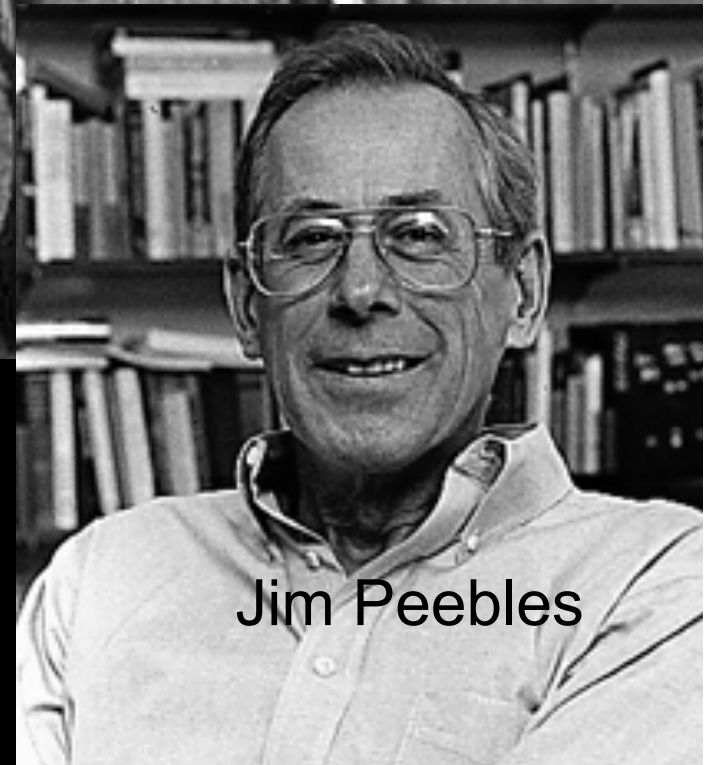
Isaac Newton



James Jeans



Yaakov Zeldovich



Jim Peebles

if the matter was evenly disposed throughout an infinite space, it could never convene into one mass; but some of it would convene into one mass and some into another, so as to make an infinite number of great masses, scattered at great distances from one to another throughout all that infinite space. And thus might the sun and fixed stars be formed, supposing the matter were of a lucid nature.

if the sun at rest were an opaque body like the planets or the planets lucid bodies like the sun, how he alone should be changed into a shining body whilst all they continue opaque, or all they be changed into opaque ones whilst he remains unchanged, I do not think explicable by mere natural causes, but am forced to ascribe it to the counsel and contrivance of a voluntary Agent.

Isaac Newton, letter to Richard Bentley, December 10, 1692

James Jeans:

"We have found that as Newton first conjectured....

All celestial bodies originate by a process of fragmentation of nebulae out of chaos, of stars out of nebulae, of planets out of stars and satellites out of planets."

Criterion for gravitational stability found by Jeans (1902):

pressure opposes collapse:

sound waves must cross region to communicate pressure changes before collapse

ASSUMPTIONS

FRW metric

Friedmann-Lemaître cosmology

primordial density fluctuations

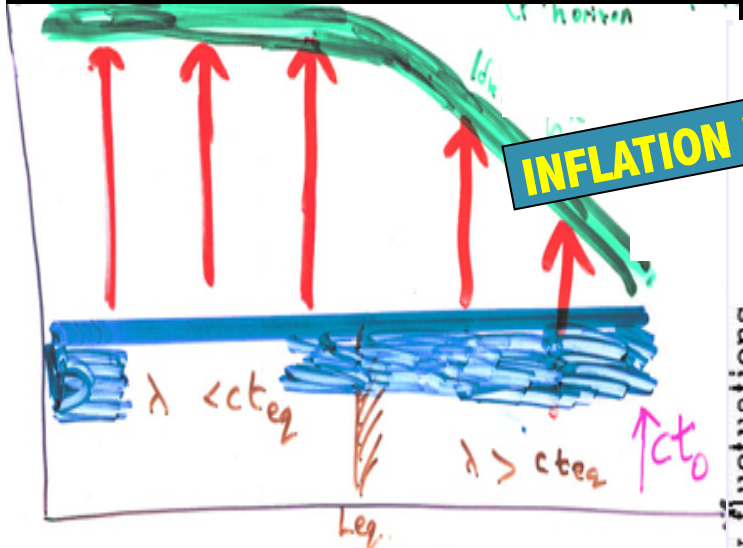
are

gaussian

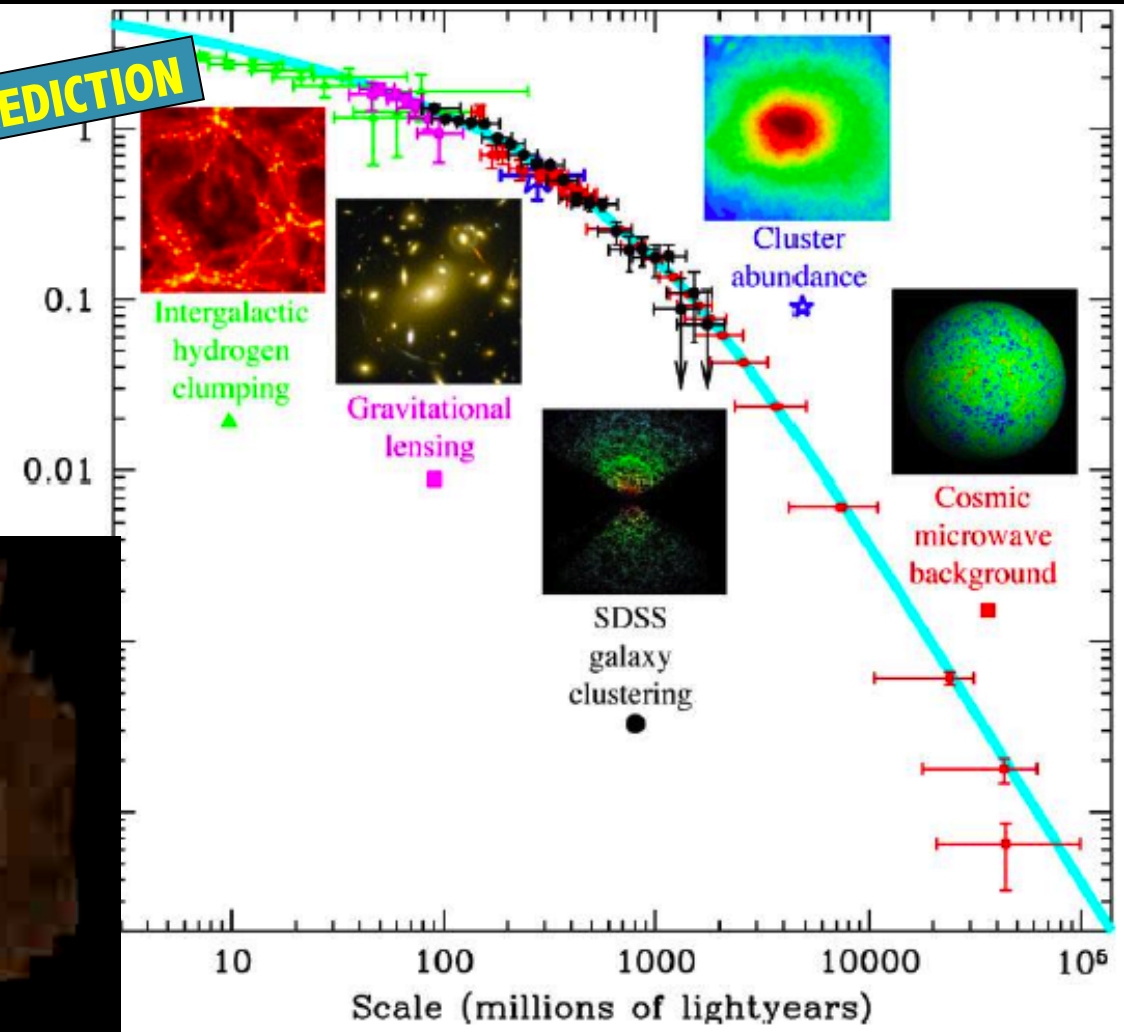
nearly scale-invariant metric
perturbations

adiabatic mode

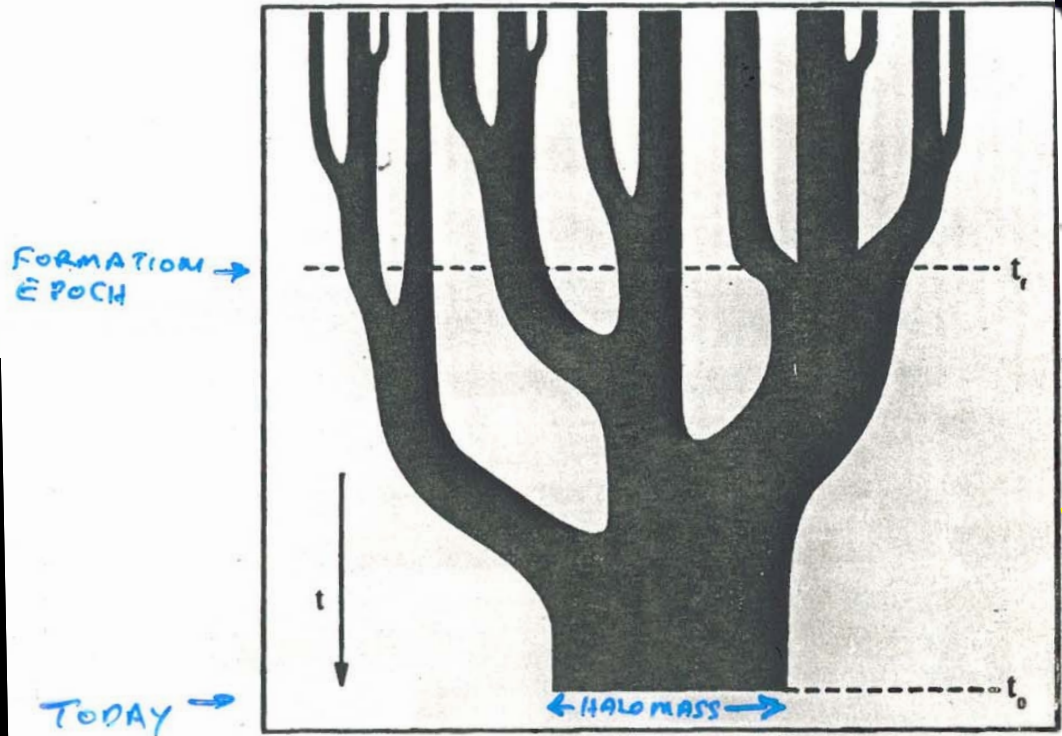
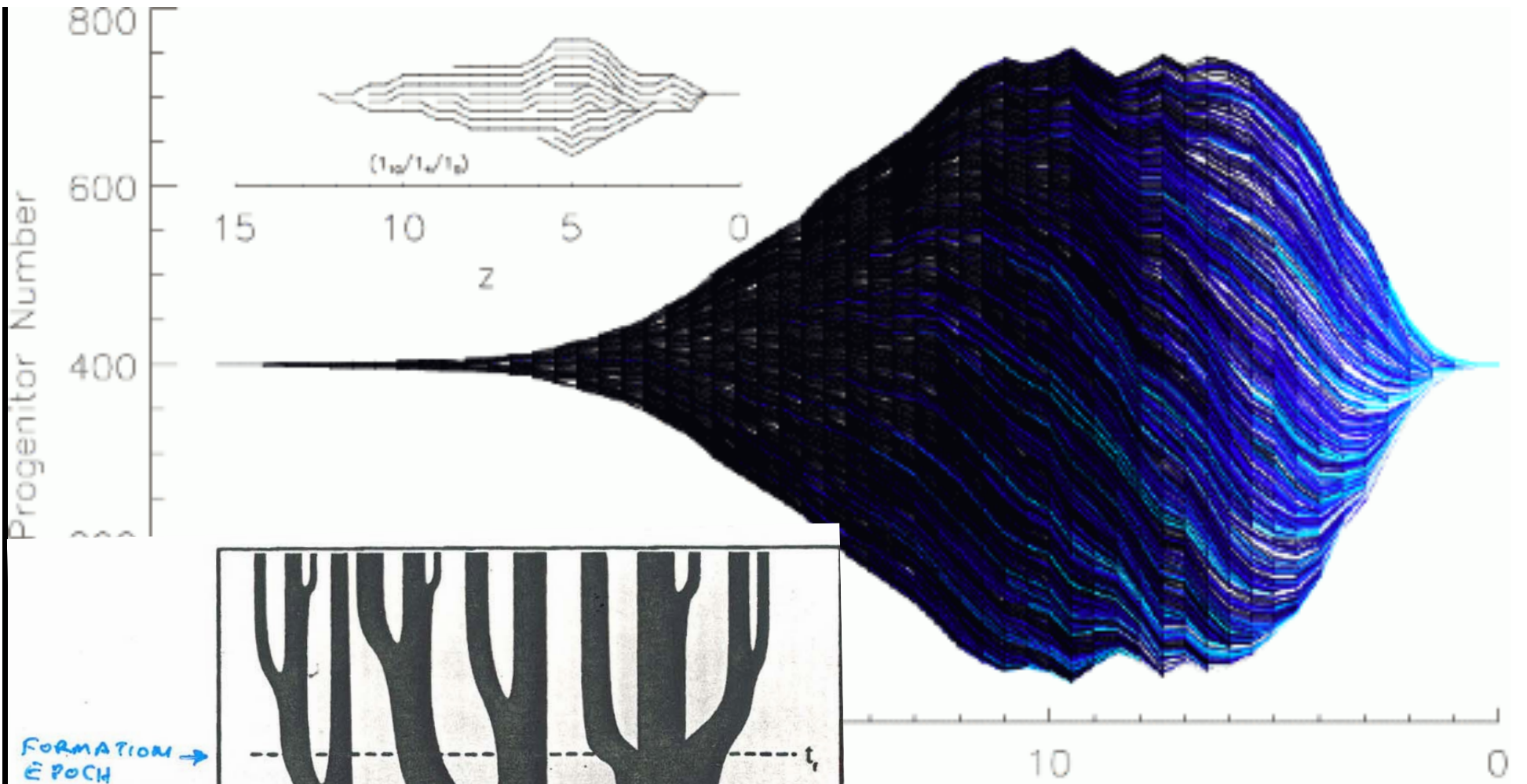
FROM INFLATIONARY DENSITY FLUCTUATIONS TO GALAXIES



INFLATION PREDICTION

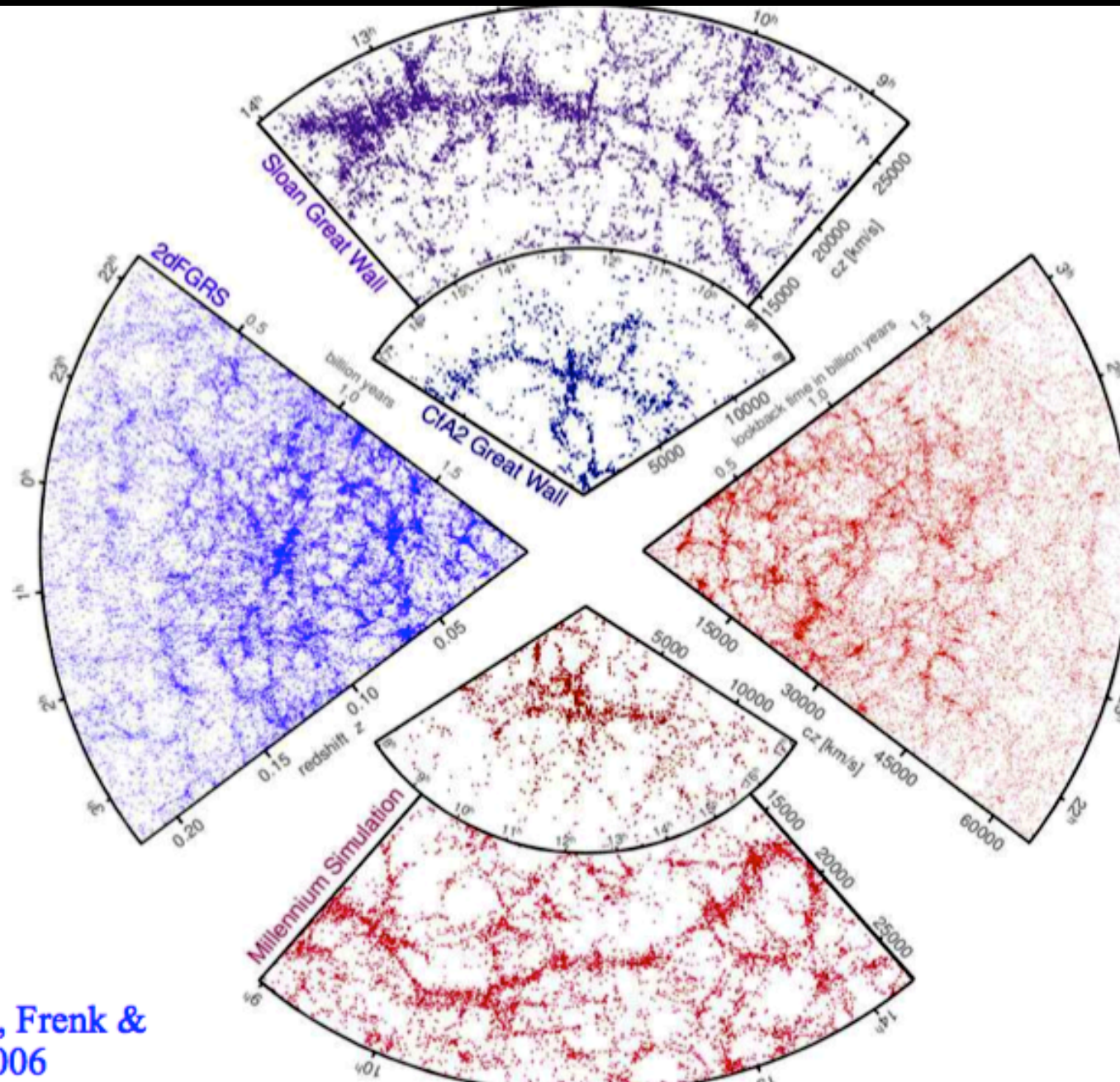


Tegmark 2004



Structure forms bottom-up

virtual universe versus observed universe



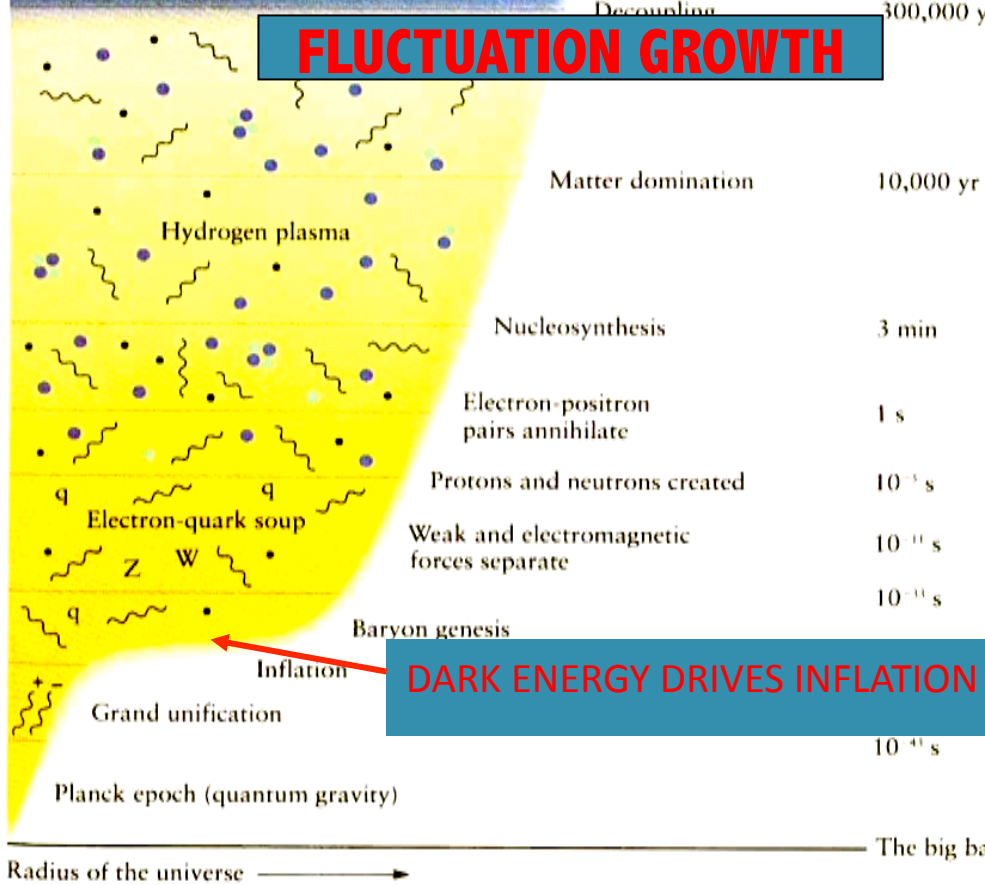
Springel, Frenk &
White 2006

A CHRONOLOGY OF THE UNIVERSE

DARK MATTER DOMINATION

Quasars	3 billion yr
Galaxy spheroids	
Protogalaxies; first stars	1 billion yr
Decoupling	300,000 yr

FLUCTUATION GROWTH



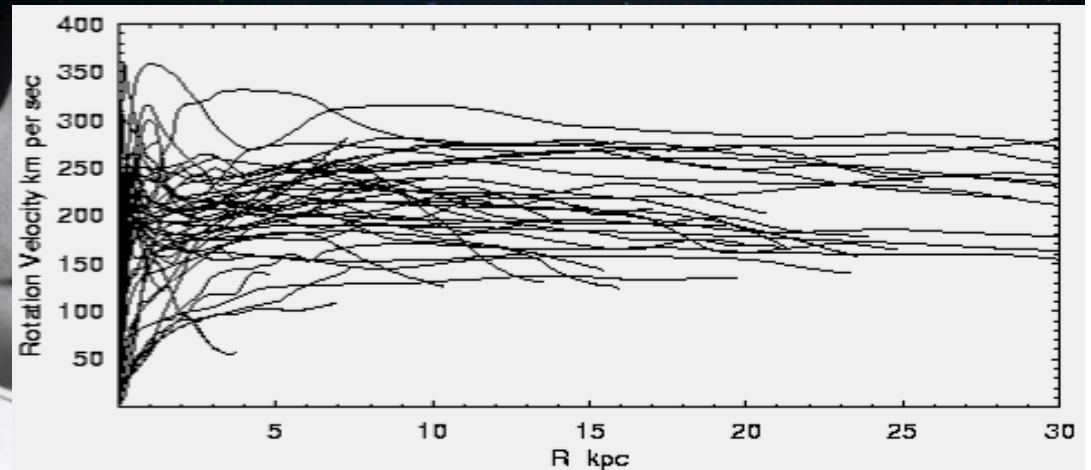
DARK ENERGY DRIVES INFLATION

Dark Matter in Galaxies

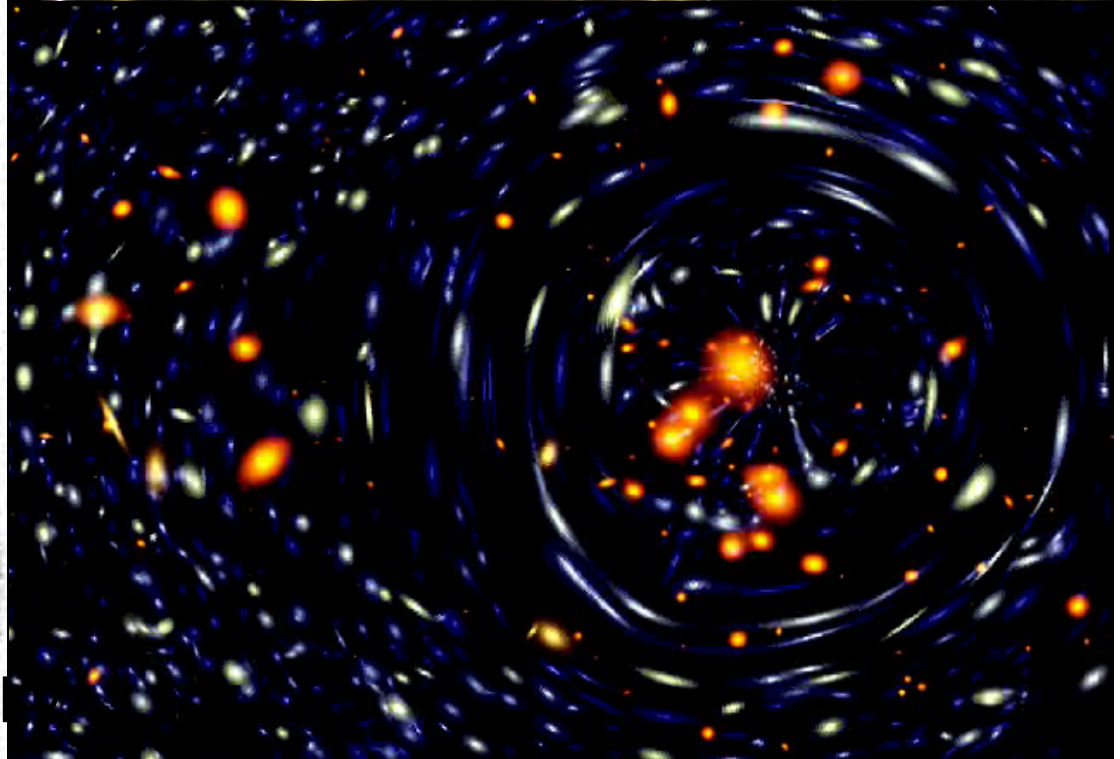
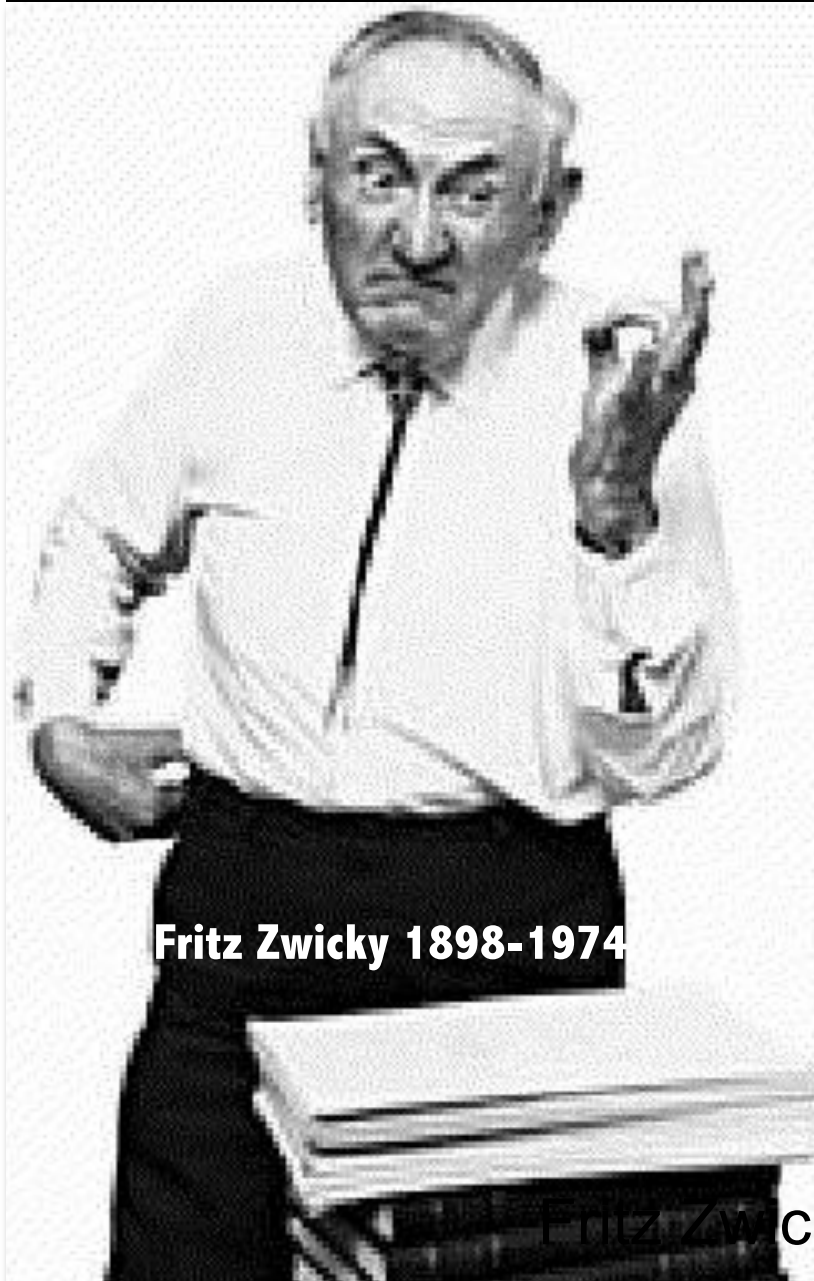
Vera Rubin 1928-



Vera Rubin

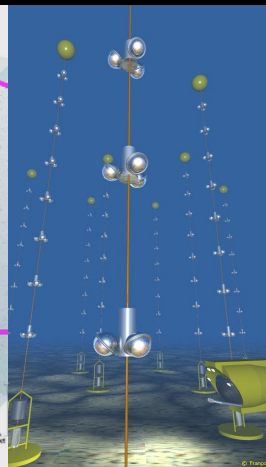
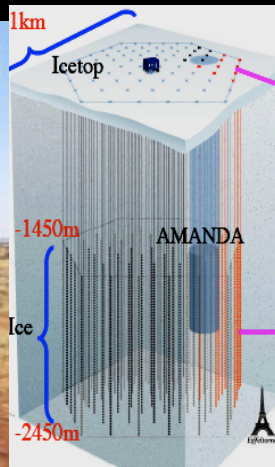


Dark Matter in Galaxy Clusters



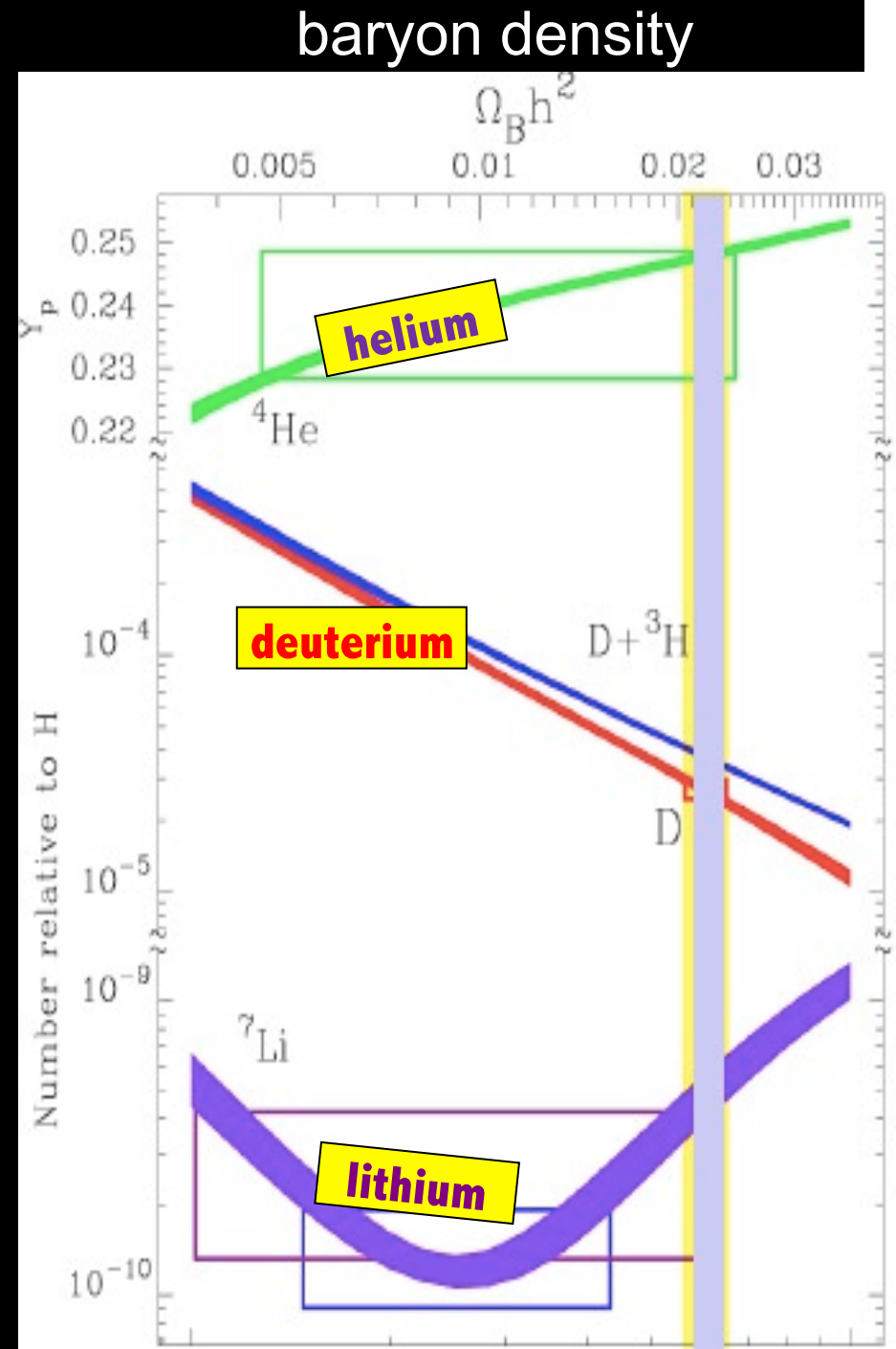
Most of the matter in the universe is dark!

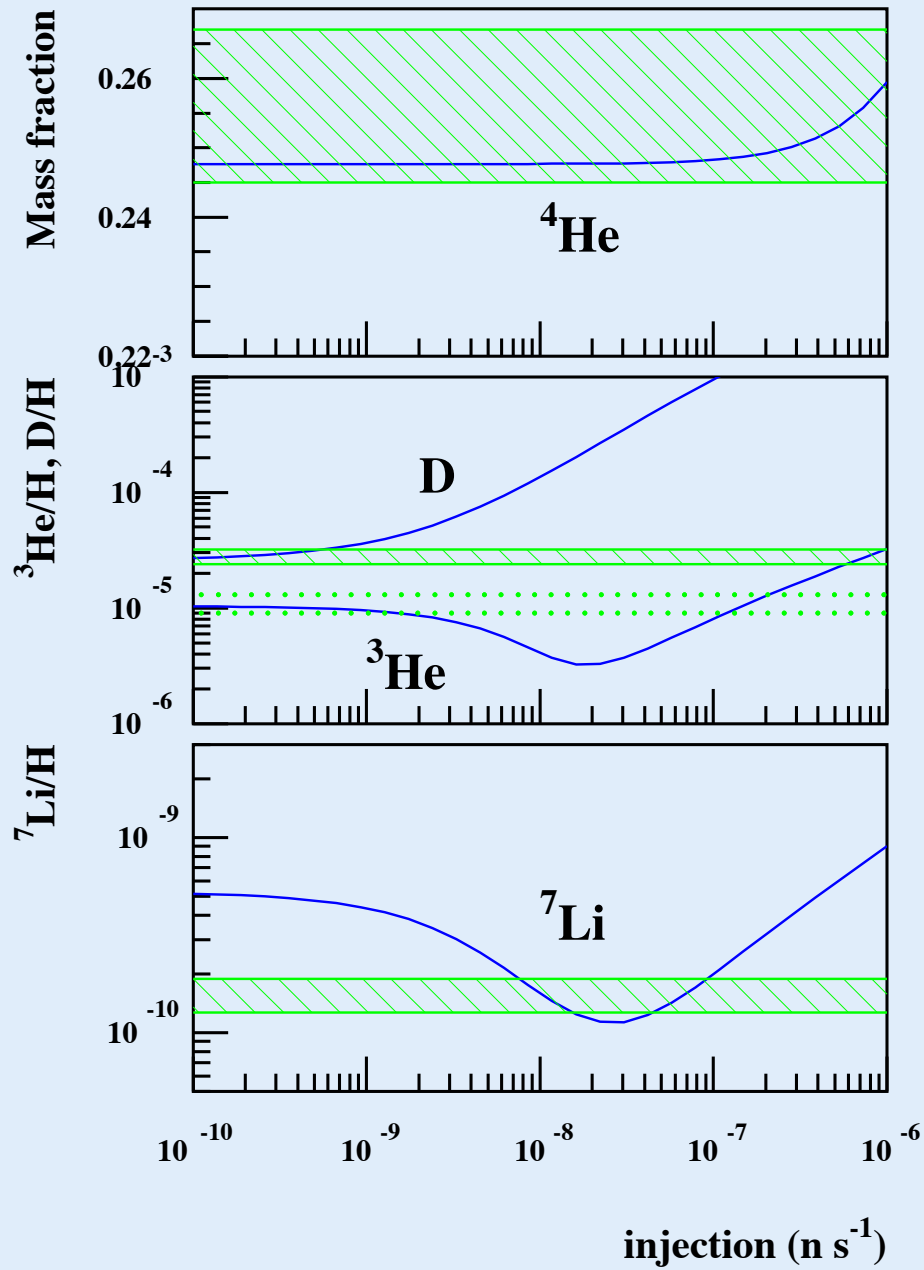
Dark matter is not ordinary matter
It's a weakly interacting elementary particle that we have yet to identify. This has led to a new discipline: Particle Astrophysics



Dark matter is not baryons

abundance





Injecting neutrons
destroys lithium
but at a price

primordial neutrinos as hot dark matter

$$\Omega_\nu h^2 = \sum m_\nu / 92 \text{ eV}$$

Hubble parameter $h = 0.65$ (65 km/s/Mpc)

$$\Omega_\nu < 0.20$$

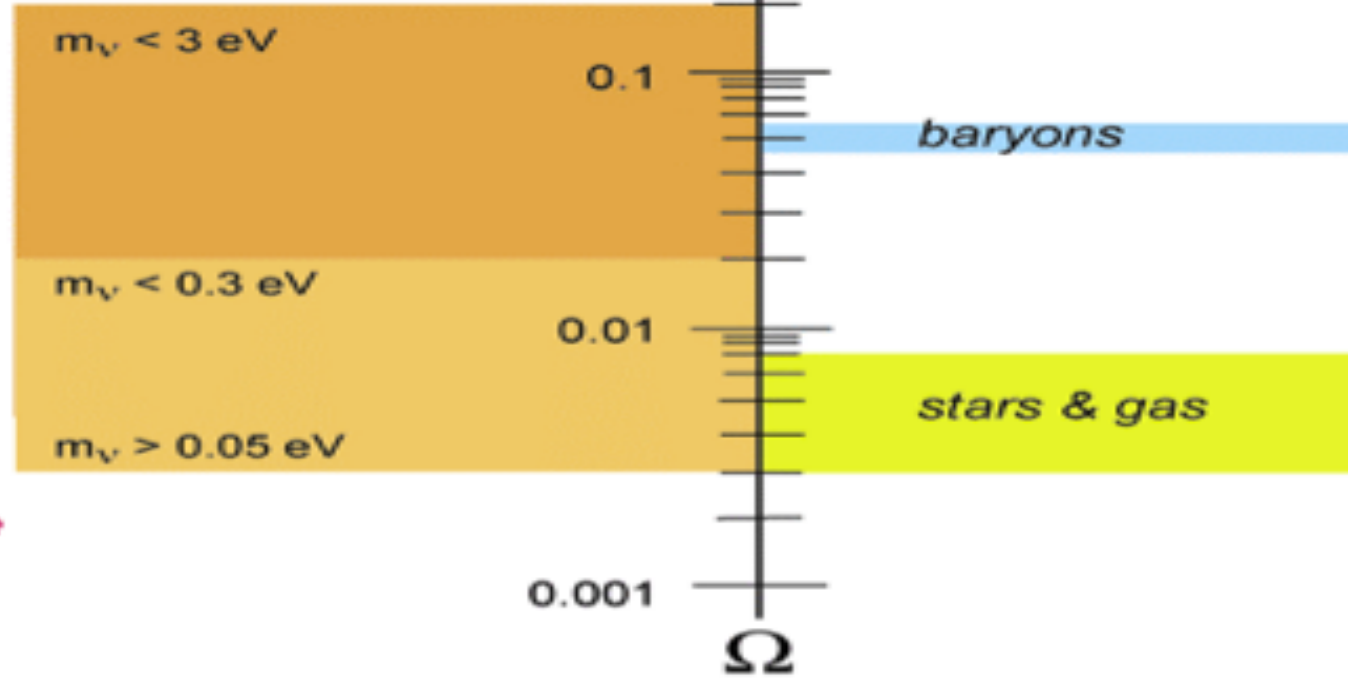
*structure formation
tritium experiments*

$$\Omega_\nu < 0.02$$

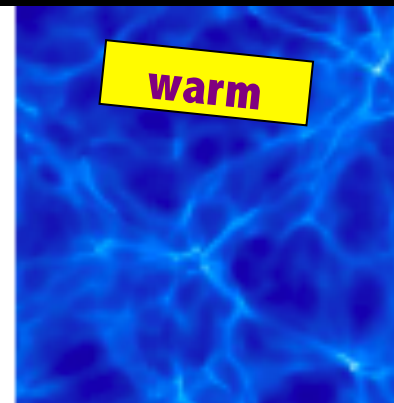
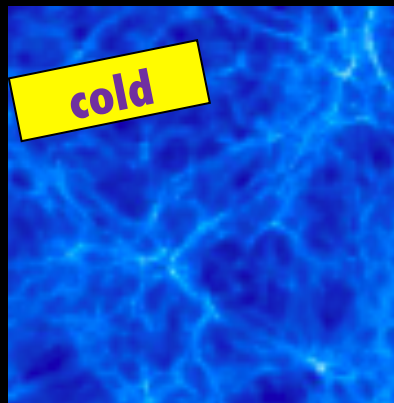
KATRIN sensitivity

$$\Omega_\nu > 0.003$$

Super-Kamiokande



Dark matter is not neutrinos

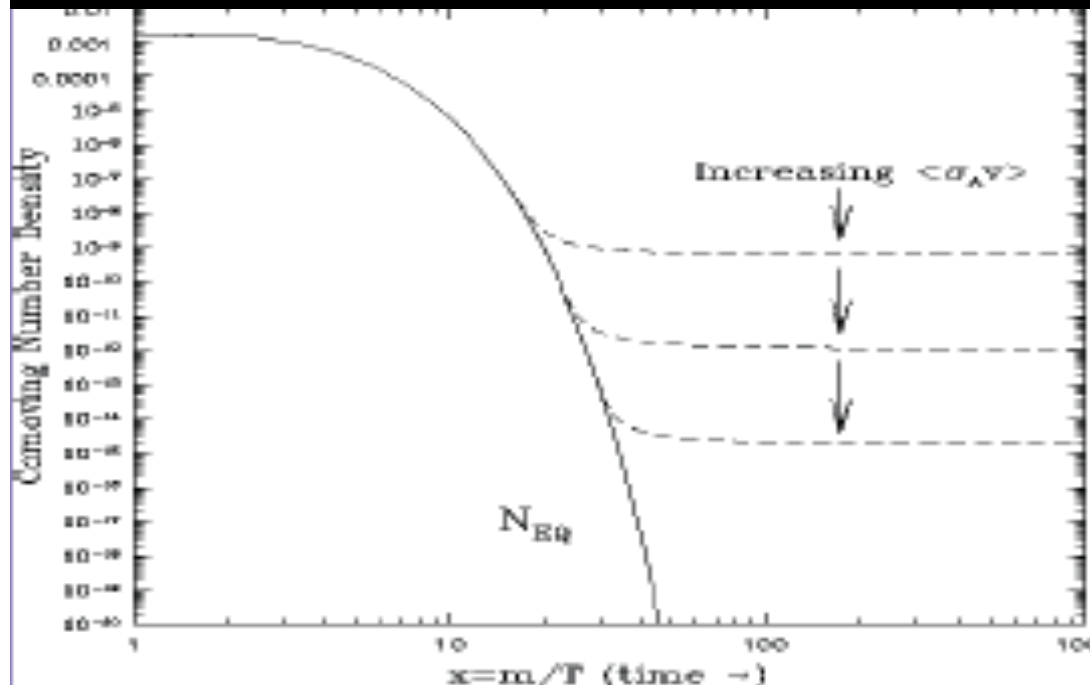


Dark matter “most likely” is
a weakly interacting (massive?) particle

Eg WIMP (or LSP) motivated by theory of supersymmetry

Favoured SUSY candidate is a WIMP in mass range 0.01-10 TeV

The WIMP miracle: relic abundance if $\langle\sigma v\rangle\sim 3\times 10^{-26}\text{ cm}^3/\text{s}\sim 1/\Omega_x$



Occasionally annihilate today
In the halo to energetic
particles: $\nu, \gamma, p\bar{p}, e^+$

Astrophysical probes complement collider experiments

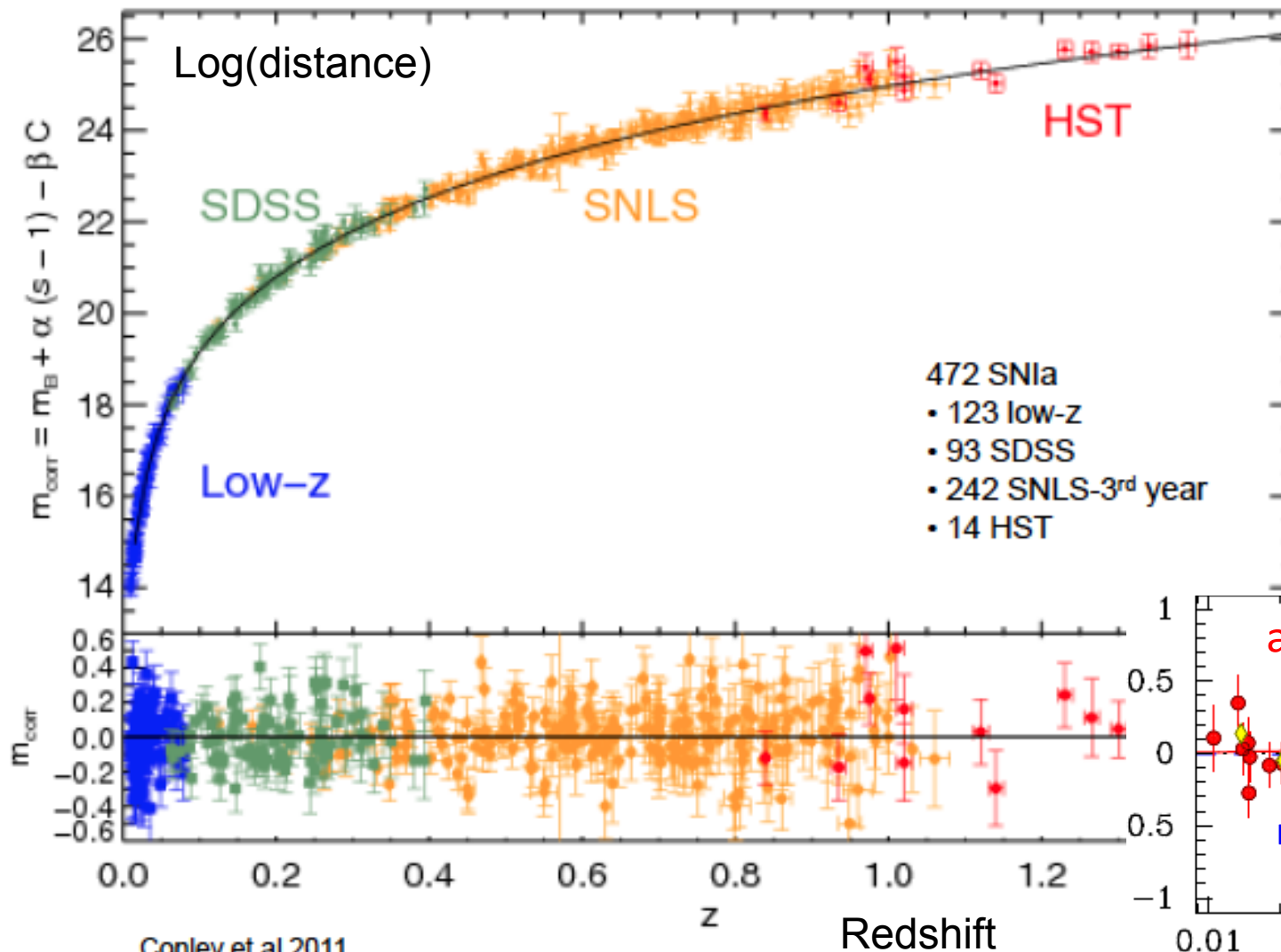
Dark energy accelerates...

we measure dark matter minus dark energy

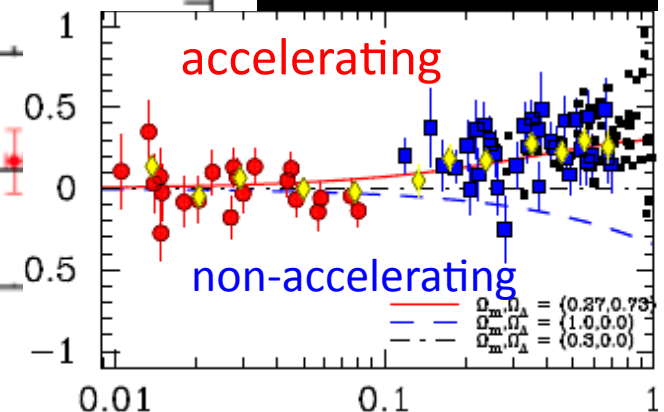
A photograph of a galaxy, likely a spiral galaxy, viewed from an angle. The galaxy's core is bright and glowing, with a reddish-pink hue. The spiral arms are visible, though somewhat faint. In the foreground, there is a bright, blue-white star with a four-pointed diffraction pattern. The background is a dark, deep blue sky with some faint, distant stars.

Distant type Ia supernovae are too faint!

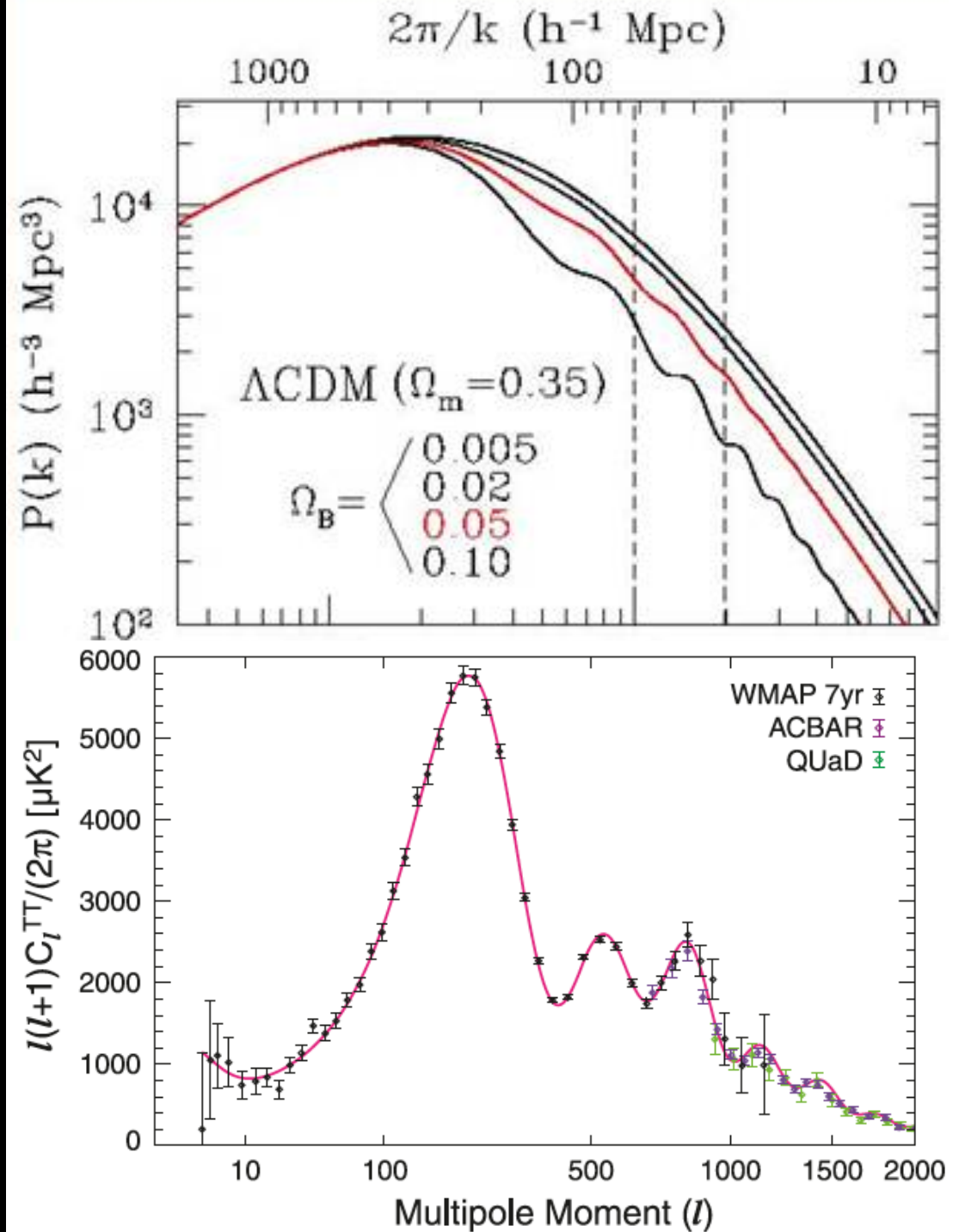
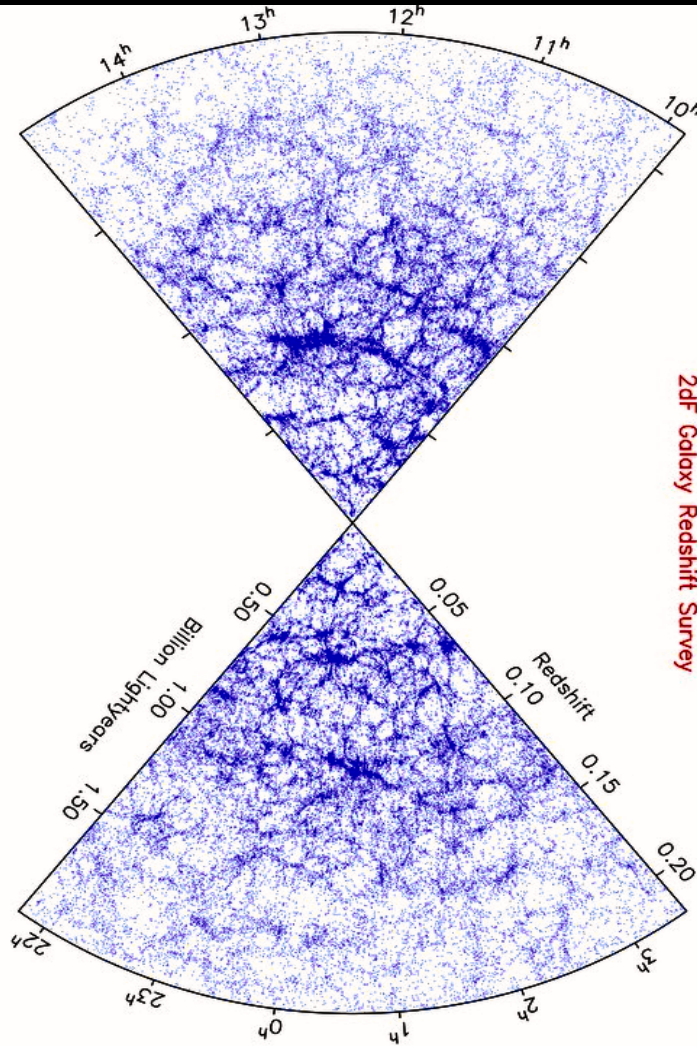
Hubble diagram SNIa

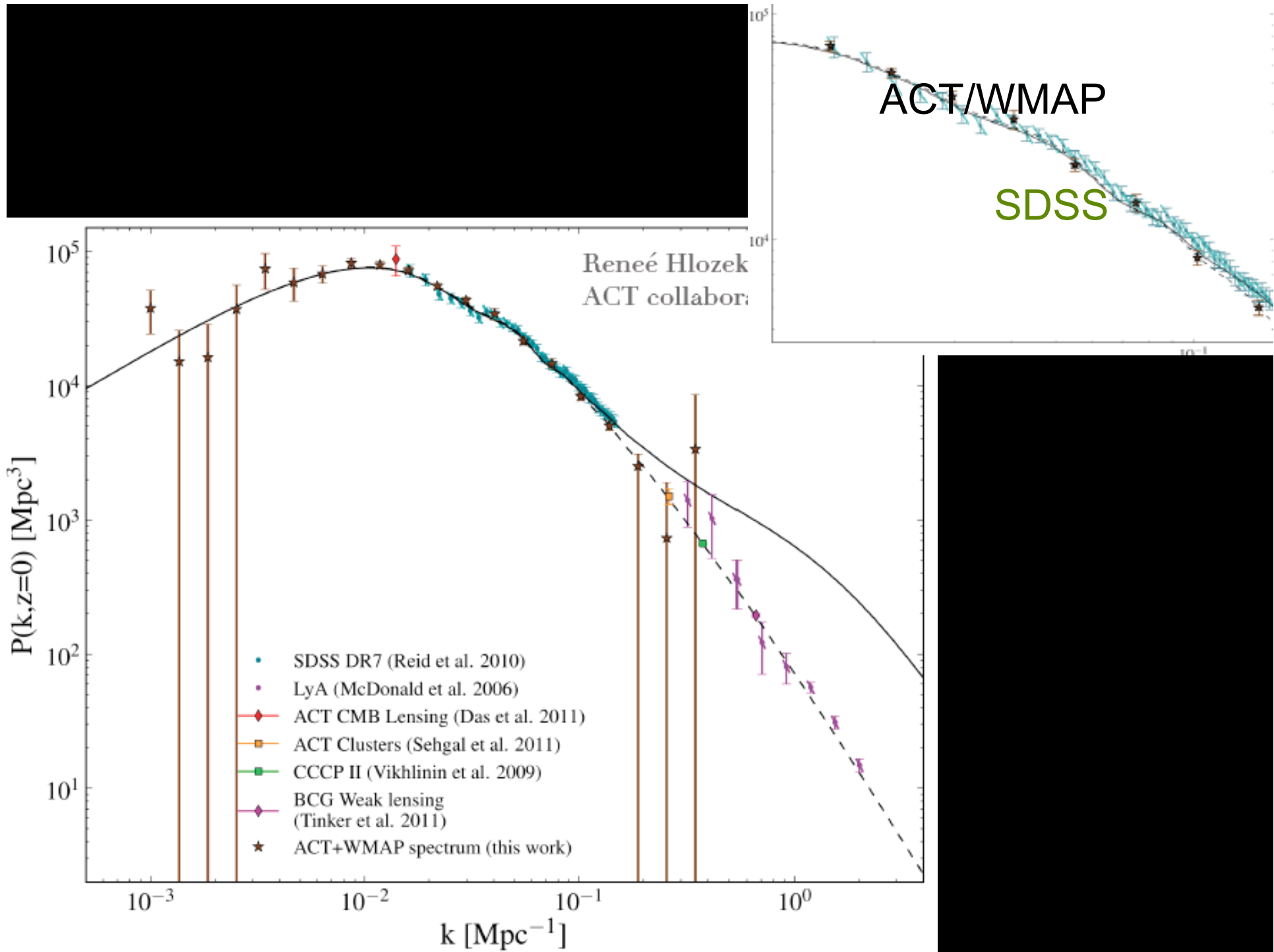


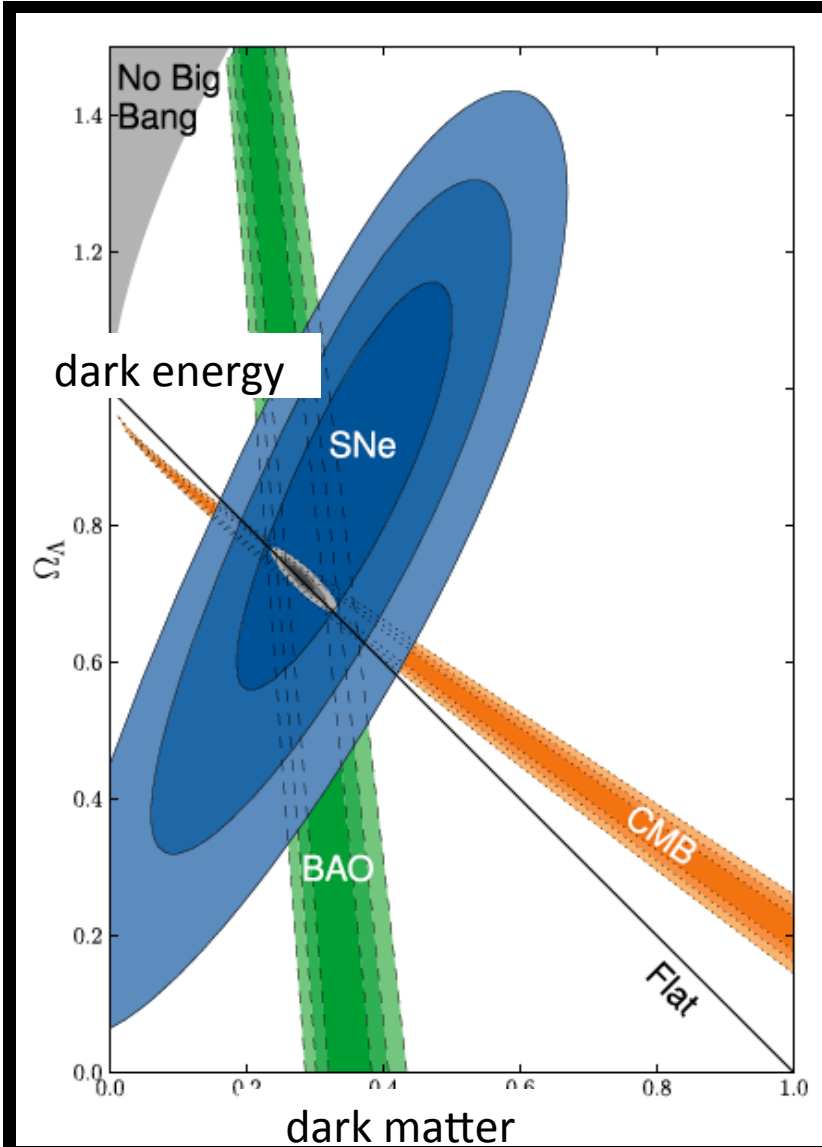
Freedman et al. 2009



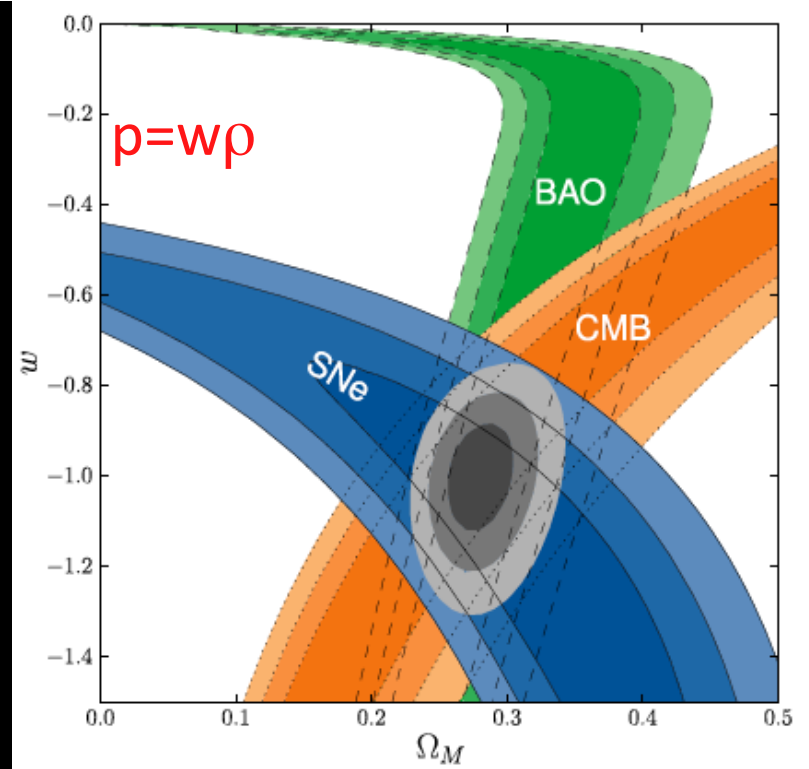
Dark matter distorts redshift space







Amanullah et al 2010



dark matter

Large-scale structure (BAO): DM

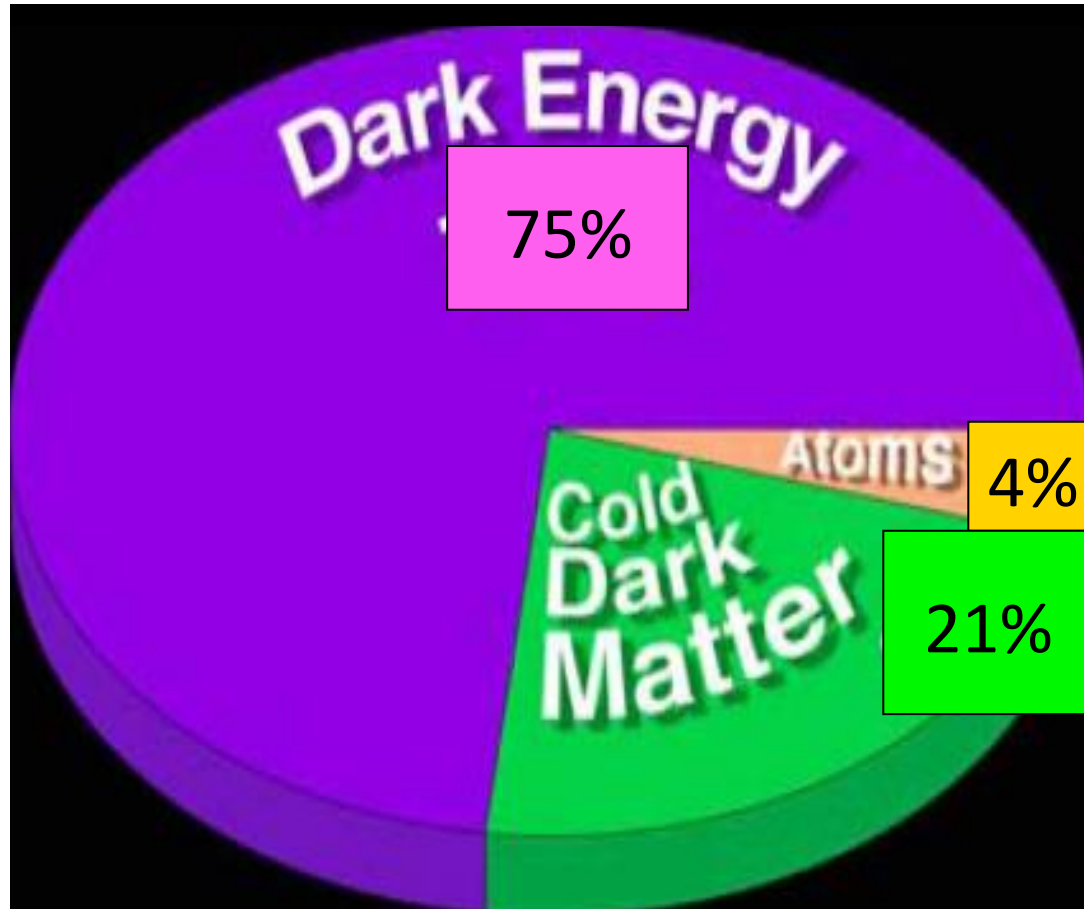
CMB fluctuations: DM + DE

Supernovae: DM - DE

Most of the energy in the universe is dark!

observe

$$\rho_{\text{vac}} \approx 10^{-10} \text{eV}^4$$



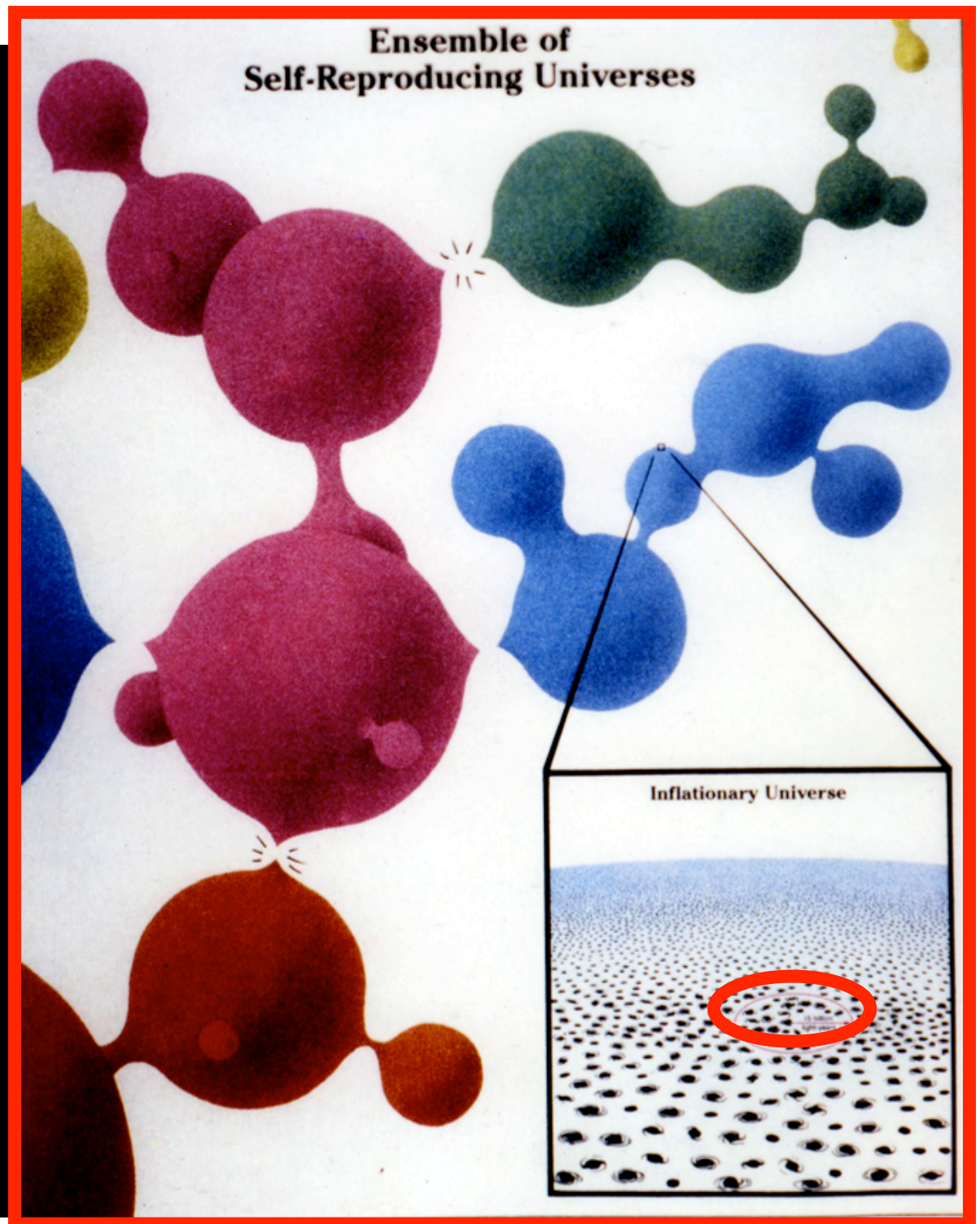
predict

$$M \sim M_{\text{Planck}} = G^{-1/2} = 10^{28} \text{eV} \Rightarrow \rho_{\text{vac}} \sim 10^{112} \text{eV}^4$$

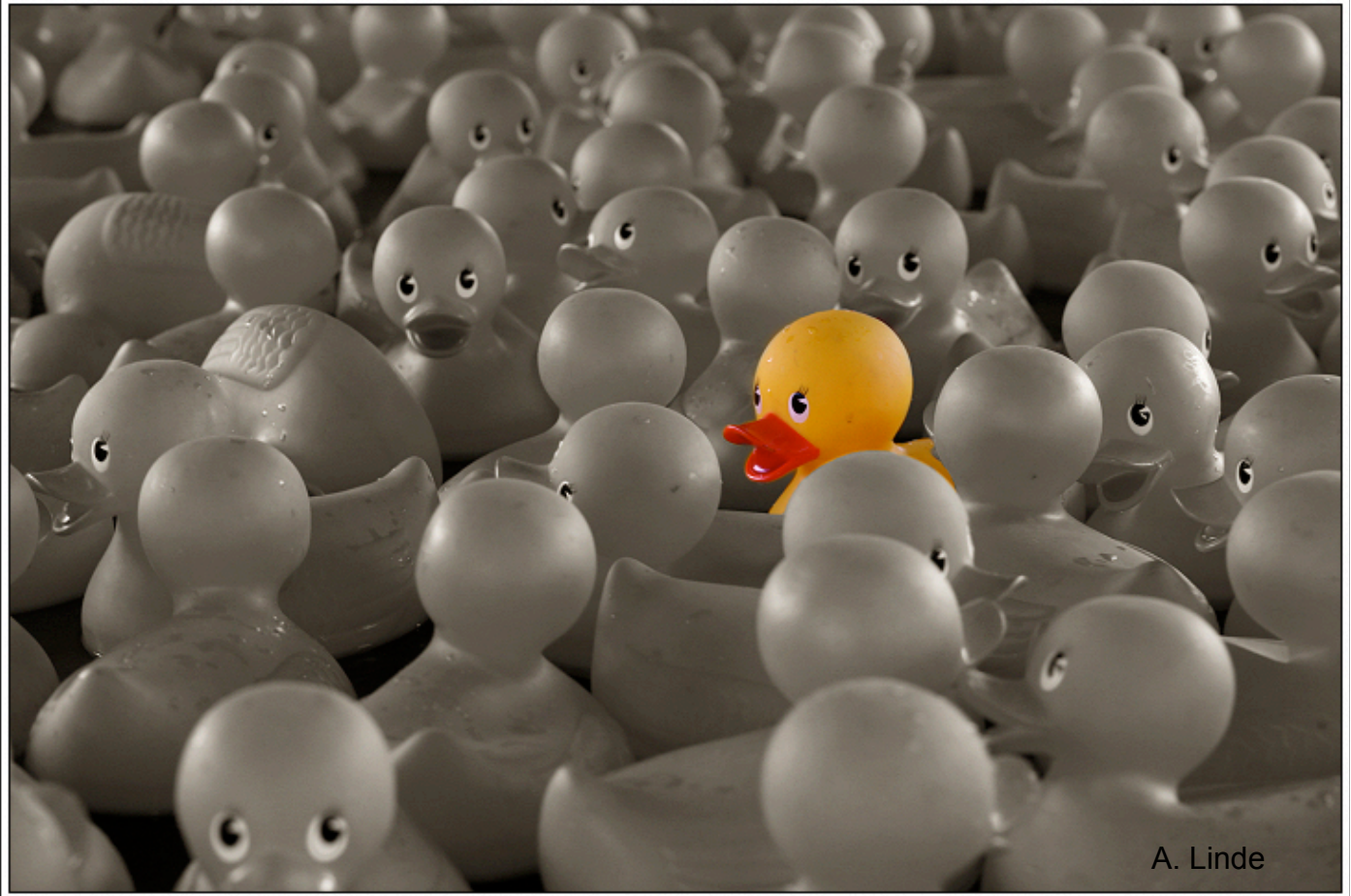
One of the greatest problems in physics!

Eternal inflation produces
a nearly infinite number of
inaccessible universes

Motivates
multiverse
explanation
of why dark
energy is so
small



We live in one tiny pocket where the value of the cosmological constant is consistent with our kind of life Leonard Susskind



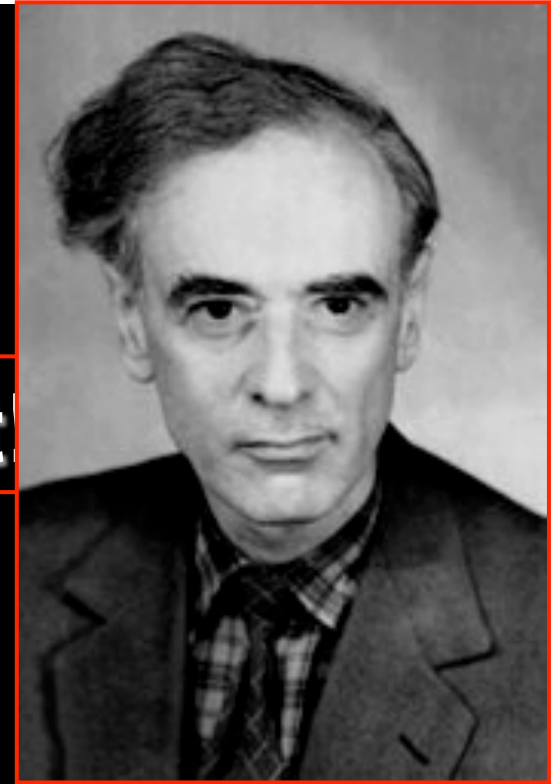
A. Linde



**The multiverse theory can't make any predictions
... it can explain anything...**
George Ellis

Landau on Cosmologists

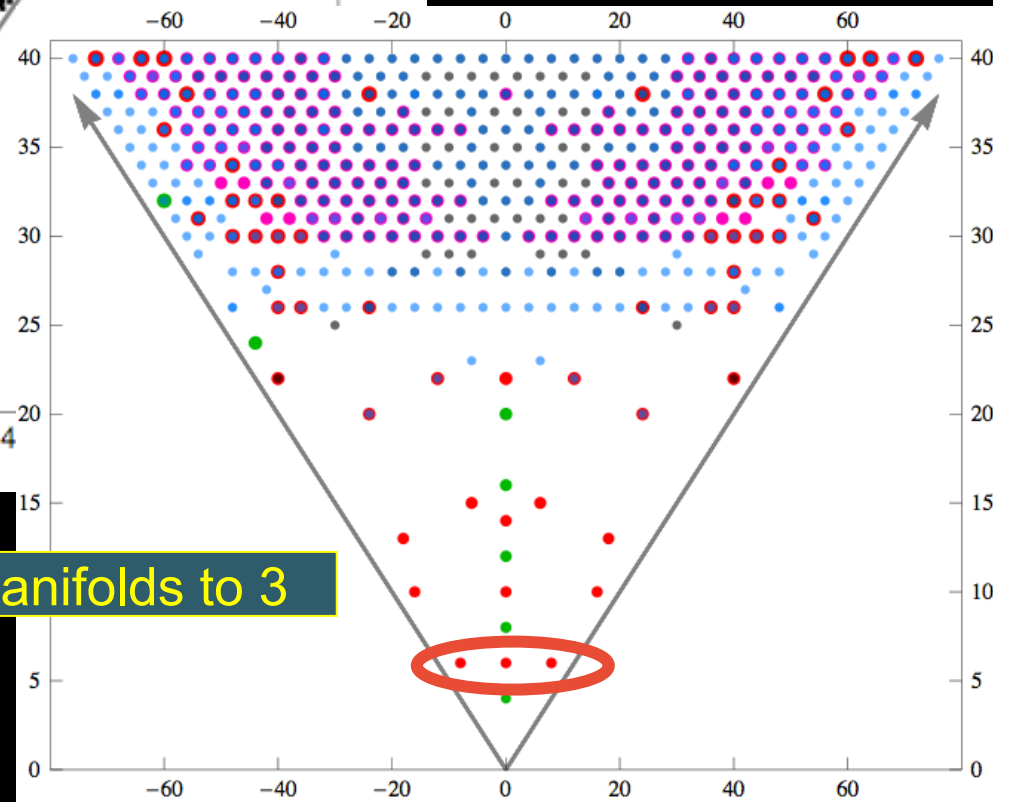
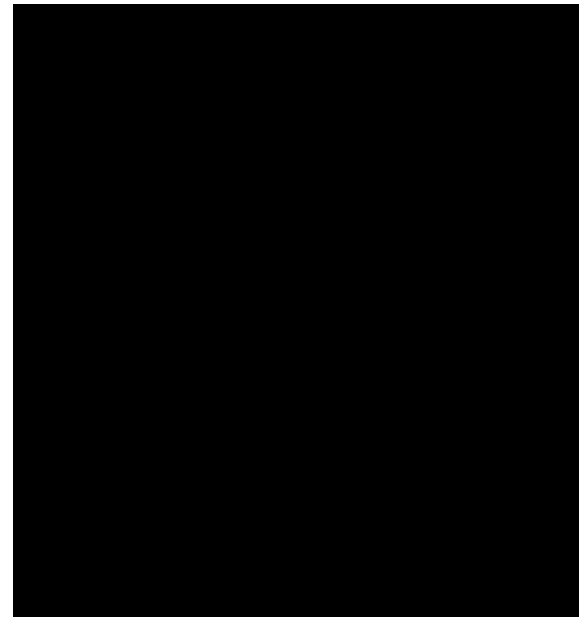
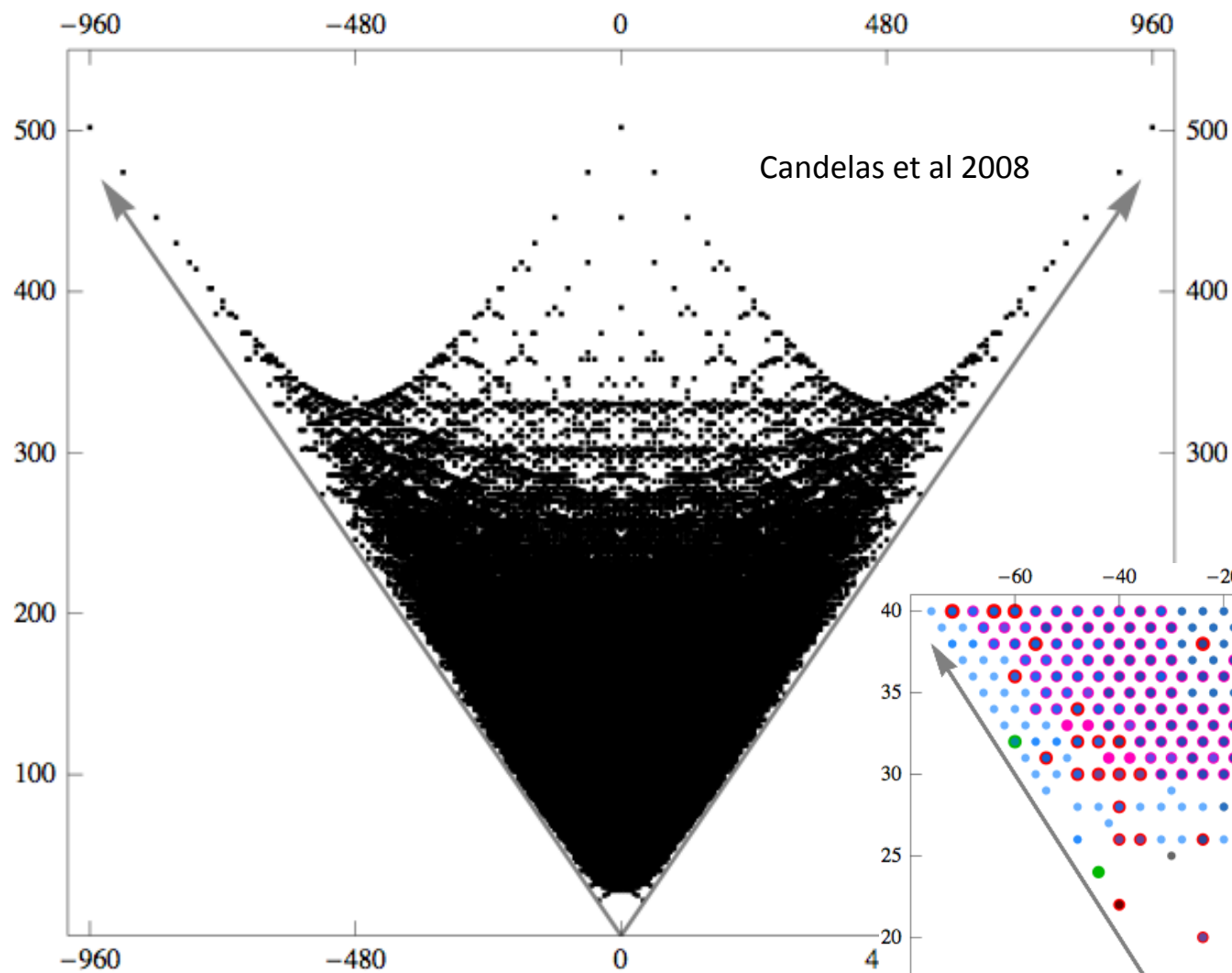
Often in Error, Never in Doubt



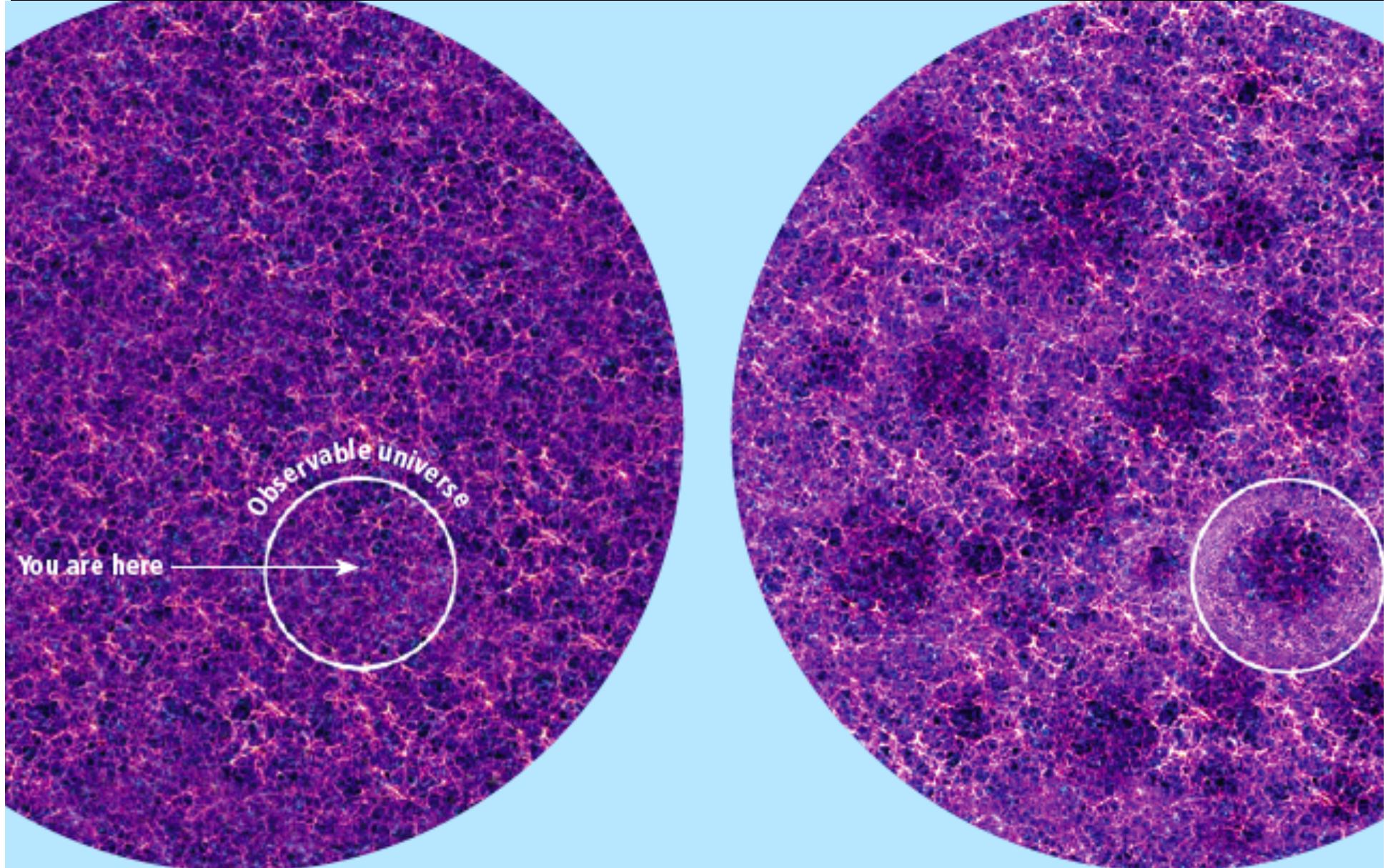
Alternatives:

we may hope for a fundamental
physics theory of dark energy.....

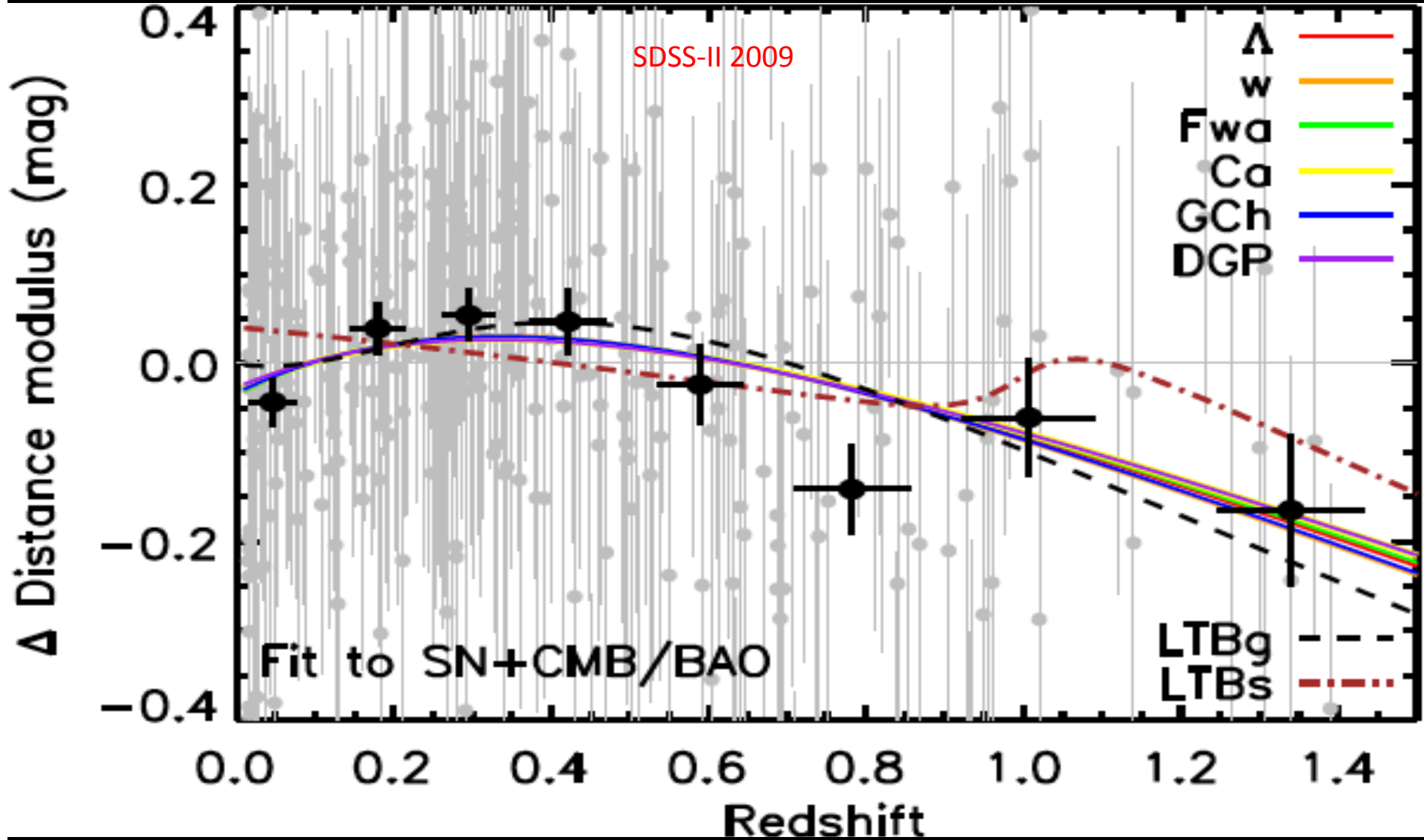
or seek an astrophysical explanation



an example: from 10^{500} Calabi-Yau manifolds to 3

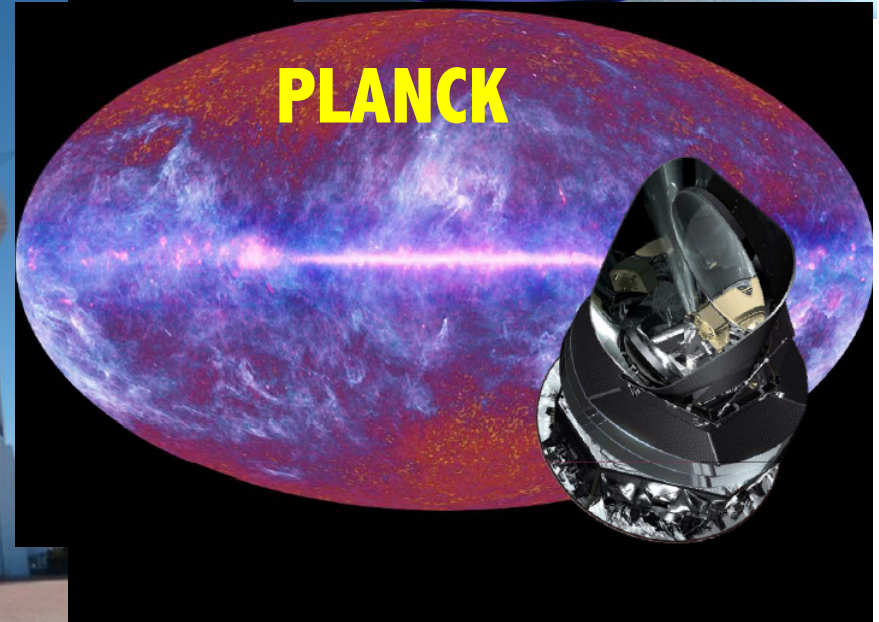
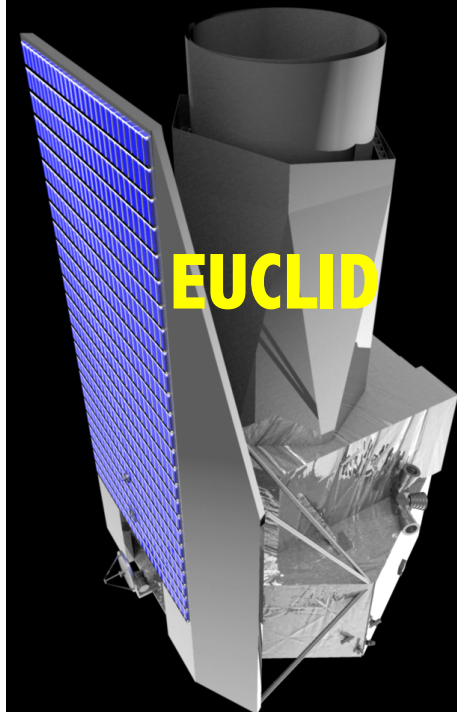
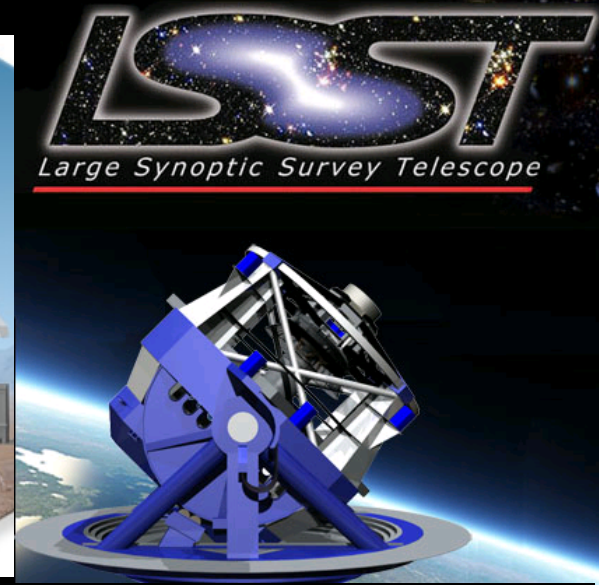


CAN REPLACE ACCELERATION BY A GIANT LOCAL VOID
in Lemaitre-Tolman-Bondi model



Which hypothesis is more “fine-tuned”?

THE FUTURE FOR DARK ENERGY



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